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Published in:
Monthly Notices of the Royal Astronomical Society

DOI:
[10.1093/mnras/stv1372](https://doi.org/10.1093/mnras/stv1372)

IMPORTANT NOTE: You are advised to consult the publisher's version (publisher's PDF) if you wish to cite from it. Please check the document version below.

Document Version
Publisher's PDF, also known as Version of record

Publication date:
2015

[Link to publication in University of Groningen/UMCG research database](#)

Citation for published version (APA):

Hess, K. M., Jarrett, T. H., Carignan, C., Passmoor, S. S., & Goedhart, S. (2015). KAT-7 science verification: Cold gas, star formation, and substructure in the nearby Antlia Cluster. *Monthly Notices of the Royal Astronomical Society*, 452(2), 1617-1636. <https://doi.org/10.1093/mnras/stv1372>

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KAT-7 science verification: cold gas, star formation, and substructure in the nearby Antlia Cluster

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Accepted 2015 June 17. Received 2015 June 17; in original form 2015 January 26

ABSTRACT

The Antlia Cluster is a nearby, dynamically young structure, and its proximity provides a valuable opportunity for detailed study of galaxy and group accretion on to clusters. We present a deep H I mosaic completed as part of spectral line commissioning of the Karoo Array Telescope (KAT-7), and identify infrared counterparts from the *Widefield Infrared Survey Explorer* extended source catalogue to study neutral atomic gas content and star formation within the cluster. We detect 37 cluster members out to a radius of ~ 0.9 Mpc with $M_{\text{H I}} > 5 \times 10^7 M_{\odot}$. Of these, 35 are new H I detections, 27 do not have previous spectroscopic redshift measurements, and one is the Compton thick Seyfert II, NGC 3281, which we detect in H I absorption. The H I galaxies lie beyond the X-ray-emitting region 200 kpc from the cluster centre and have experienced ram pressure stripping out to at least 600 kpc. At larger radii, they are distributed asymmetrically suggesting accretion from surrounding filaments. Combining H I with optical redshifts, we perform a detailed dynamical analysis of the internal substructure, identify large infalling groups, and present the first compilation of the large-scale distribution of H I and star-forming galaxies within the cluster. We find that elliptical galaxy NGC 3268 is at the centre of the oldest substructure and argue that NGC 3258 and its companion population are more recent arrivals. Through the presence of H I and ongoing star formation, we rank substructures with respect to their relative time since accretion on to Antlia.

Key words: galaxies: clusters: individual: Antlia Cluster – galaxies: distances and redshifts – galaxies: evolution – galaxies: kinematics and dynamics – infrared: galaxies – radio lines: galaxies.

1 INTRODUCTION

Galaxy clusters are the product of collapse, assembly, and merging of dark matter haloes over cosmic time in a Λ cold dark matter universe (Springel et al. 2005), and we know that hierarchical structure formation is at work in early times as galaxy clusters are well established by at least $z = 1$ (e.g. Fassbender et al. 2011; Gettings et al. 2012, and references therein). Galaxies may be accreted on to clusters individually or as members of galaxy groups.

On the outskirts of clusters, galaxy transformation occurs on relatively rapid time-scales, while within the cluster, dynamical friction smoothes out galaxy distributions, so the evidence of cluster assembly decays with time. None the less, nearly all clusters display an element of youth through the presence of substructure: a signature of their merger history. Understanding the assembly history of clusters through cosmic time is important for both understanding the evolution of large-scale structure, and the properties of the galaxy population within it.

Theoretically, 30 per cent of all systems should contain substructure, although observations report a range of 27–73 per cent for massive groups (more than 20 members) and clusters (Dressler &

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Shectman 1988; Bird 1994; Ramella et al. 2007; Hou et al. 2012). Substructure tends to be anti-biased with respect to the dark matter distribution, preferentially surviving on the outskirts of systems (Oguri & Lee 2004), but it is unclear how long infallen substructure may survive as both intermediate- and low-redshift clusters can host density enhancements and multiple massive galaxies (e.g. Beers & Geller 1983). Even in the interior of clusters, cD galaxies typically reside at the centre of galaxy density enhancements rather than at the centre of the clusters (Zabludoff, Huchra & Geller 1990; Beers et al. 1991; Bird 1994), suggesting themselves originated from the merger of groups (Zabludoff & Mulchaey 1998).

In general, substructure is identified through line-of-sight velocities, and the spatial distribution of small collections of galaxies relative to the overall cluster population (Dressler & Shectman 1988; Bird 1994; Serna & Gerbal 1996; Jaffé et al. 2012; Pranger et al. 2013). Despite a bevy of statistical tools at our disposal, we are still striving to understand to what degree clusters are dynamically relaxed or are continuing to accrete from their environment, even at low redshift (Hou, Parker & Harris 2014).

A spectroscopic census of the stellar content of clusters illuminates these substructures, however the presence of atomic gas, measured directly by H I 21 cm observations, not only provides a complimentary evidence of substructure (Bravo-Alfaro et al. 2000; Jaffé et al. 2012), but also indicates which galaxies have most recently accreted from the surrounding environment. Late-type galaxies in clusters are H I deficient with respect to their counterparts in the field (e.g. Haynes & Giovanelli 1984; Solanes et al. 2001). They have shrunken H I discs where X-ray-emitting hot intracluster gas fills the cluster potential well (Bravo-Alfaro et al. 2000; Chung et al. 2009) and high-resolution H I images of Virgo cluster members show that the hot intracluster medium is responsible in many cases for removing the gas through ram pressure stripping (Kenney, van Gorkom & Vollmer 2004; Cowl et al. 2005; Chung et al. 2009; Jaffé et al. 2015). In most cases, H I objects avoid the centres of clusters all together.

The removal of the atomic gas, in the long run, leads to the shut down of star formation and a colour transformation (Schawinski et al. 2014), and environment is responsible for quenching star formation in galaxies across a range of parent dark matter halo masses (e.g. Wetzel, Tinker & Conroy 2012; Woo et al. 2013). However, the impact of the cluster environment on star formation within an individual galaxy is not straight forward: the *fraction* of star-forming galaxies in clusters increases with radius until reaching the field value at 2–3 virial radii, but the mean star formation *rate* shows no radial dependence (Rines et al. 2005). Similarly, despite cluster galaxies being H I deficient, CO, which traces the cold H₂ out of which stars form, shows no environmental dependence (Kenney & Young 1989; Casoli et al. 1991). Thus, quenching may be delayed if galaxies have a sizeable reservoir of molecular gas. These complications persist despite the fact that the distance at which a galaxy resides from the cluster centre is correlated with the time since infall (De Lucia et al. 2012). Combining H I and star formation proxies, a cluster member which has star formation, but is not detected in H I may indicate an intermediate age cluster member, stripped of its H I but converting molecular gas to stars.

The infall region of clusters is an interesting volume in which to study galaxy evolution because the rapid change in galaxy density and properties of the intergalactic medium between the field and cluster environment have a strong impact on the stars and interstellar medium (ISM). However, a number of studies show that local density is a more important factor determining the present day star formation and gas content of galaxies than large-scale

environment suggesting that ‘pre-processing’ is important in intermediate density environments before galaxies arrive in clusters (e.g. Rines et al. 2005; Lane et al. 2007; Blanton & Berlind 2007; Hess 2011). Disentangling the relative importance of the ‘pre-processing’ that happens to galaxies in groups before they are accreted (e.g. Zabludoff & Mulchaey 1998; Hess & Wilcots 2013; Berrier et al. 2009), versus the impact of the cluster environment is an ongoing challenge.

Hou et al. (2014) tackle this by combining optical group catalogues with the Dressler–Shectman (DS) statistical test (Dressler & Shectman 1988) to identify subhaloes which they classify as infall, backplash, or virialized. They find evidence for ‘enhanced quenching’ of the star formation in infalling subhaloes, which they attribute to ‘pre-processing’ in group haloes, but that it is only important for the most massive accreting systems, $>10^{14.5} M_{\odot}$. This is unlikely to be the complete picture, but has provided insight into the most massive systems, and the mergers of haloes with the highest mass ratios. Complementarily, studies of the H I content of galaxies and semi-analytical models show that environmental effects become important in groups with $>10^{13} M_{\odot}$ of dark matter (McGee et al. 2009; Jaffé et al. 2012; Hess & Wilcots 2013) and Jaffé et al. (2015) show gas-poor galaxies in the infall region of Abell 963 were likely pre-processed in their previous environment.

Estimates for the contribution of cluster stellar mass that has been accreted through galaxy groups (and the redshift at which this is important) range from 12 to 50 per cent (Berrier et al. 2009; McGee et al. 2009; De Lucia et al. 2012). When galaxies fall into clusters as groups, pre-processing which has already occurred in the group environment likely influences the rate at which quenching or morphological transformation occurs. However, if galaxies fall in as individuals from the field, the interaction with the intracluster medium may induce quenching on more rapid time-scales. Understanding these different evolutionary paths may help us further understand the resulting cluster populations and the morphological transformation towards early-type and S0 galaxies (Schawinski et al. 2014).

The number of nearby galaxy clusters in which we can perform a detailed, resolved study of both the stellar and gaseous component of galaxies is limited, but include systems with a range of apparent dynamical ages from quite old such as Coma, through Virgo and Fornax, to quite young such as Abell 1367. At ~ 40 Mpc, the Antlia Cluster (Abell S0636) is the third most nearby cluster after Virgo and Fornax, yet it has been poorly studied due to its relatively low Galactic latitude ($b = 19^\circ$) and southern declination ($\delta = -35^\circ 3'$). It was not covered by the H I Parkes Zone of Avoidance Survey (Kraan-Korteweg, Henning & Schröder 2002), and only recently has there been a significant survey of the stellar content of the cluster (Smith Castelli et al. 2008, 2012). Vaduvescu et al. (2014) provide an excellent review of the optical photometric, spectroscopic, and X-ray campaigns of the Antlia Cluster, which are very incomplete.

From what we do know, Antlia is a compelling target for further study. Ferguson & Sandage (1990) found that the galaxy density is 1.7 times higher than Virgo and 1.4 times higher than Fornax. Smith Castelli et al. (2008) note that S0s outnumber elliptical galaxies by a factor of 3, contrary to Virgo and Fornax where ellipticals dominate. Together, this suggests that Antlia is the youngest of the three systems. Further, if ellipticals and S0s have different formation mechanisms, it may indicate that the clusters formed differently, if not only on different time-scales.

Overall, the cluster contains an estimated 375 galaxies (Ferguson & Sandage 1990), including at least two dozen confirmed early-type dwarfs (Smith Castelli et al. 2012). At the centre of Antlia

are two concentrations of galaxies focused around the massive ellipticals, NGC 3268 and NGC 3258. They have the same systemic velocity, but these ellipticals each host an extended globular cluster population (Dirsch, Richtler & Bassino 2003), and hot X-ray halo (Pedersen, Yoshii & Sommer-Larsen 1997; Nakazawa et al. 2000), suggesting that the galaxies are the product of past mergers, and they sit at the centre of their own subhalo. Antlia itself lacks a central X-ray excess that is typical of poor clusters, but the X-ray emission and globular cluster systems around NGC 3268 and NGC 3258 are extended along the line connecting the two galaxies, suggesting that now the two systems reside within the same dark matter halo.

Previous studies of Antlia detail the properties of the oldest merged components within the cluster: early-type galaxies at the centre embedded in hot X-ray haloes, and their early-type dwarf companions. In this work, we combine $H\text{I}$ spectral line commissioning observations from the Karoo Array Telescope (KAT-7) with infrared proxies of the stellar mass and star formation rate of galaxies from the *Widefield Infrared Survey Explorer* (WISE), and optical spectroscopic redshifts to disentangle the history of mass assembly. The $H\text{I}$ indicates recent infall to the cluster, and star formation persists even in galaxies where the gas content has been exhausted or dropped below current detection limits. A deeper understanding of Antlia, and its evolutionary state provides an important addition to the detailed knowledge of other nearby clusters.

Throughout this paper, we adopt for the cluster a heliocentric velocity of 2797 km s^{-1} ($z = 0.00933$; Smith Castelli et al. 2008). We assume a concordance cosmology with $H_0 = 73 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_M = 0.27$, $\Omega_\Lambda = 0.73$, putting Antlia at a distance of $38.1 \pm 2.7 \text{ Mpc}$, corrected for Virgo, the Great Attractor, and the Shapley supercluster.

2 KAT-7 OBSERVATIONS AND COMMISSIONING

The seven dish Karoo Array Telescope, KAT-7 (Carignan et al. 2013; Foley et al., submitted) is a test bed instrument for the South African precursor to the Square Kilometre Array (SKA), MeerKAT. We observed the Antlia Cluster with KAT-7 using a seven pointing mosaic centred on NGC 3268 (RA(2000) = $10^{\text{h}}30^{\text{m}}03^{\text{s}}.6$, Dec.(2000) = $-35^{\circ}19'23''.88$). The pointings were arranged in a discrete hexagonal pattern in the style of the NRAO VLA Sky Survey (Condon et al. 1998) with pointing centres separated by $1/\sqrt{2}$ times the half power beamwidth.¹ The total area of the mosaic covers approximately 4.4 deg^2 .

These observations were done as part of commissioning the ‘c16n25M4k’ $H\text{I}$ spectral line correlator mode of KAT-7 which features 4096 channels covering 25 MHz, with 6.104 kHz spectral resolution ($v = 1.29 \text{ km s}^{-1}$ at $z = 0$). We centred the band at 1406.7 MHz (2922 km s^{-1}) and cycled through the full mosaic with 4 minutes per pointing between 3 minutes observations of the complex gain calibrator, 1018-317. The bandpass calibrators, 3C138 and PKS 1934-638, were observed once every 2 h. The data were taken between 2013 March and 2013 July over 19 observing sessions. Sessions typically spanned 9.5–12.75 h for a total of 192 h, including 147 h on the mosaic with 21 h per pointing. Table 1 provides a summary of the observing and imaging parameters.

Table 1. KAT-7 observations and imaging.

Observing parameter	Value
Date of observations	2013 March–July
Total integration time ($7 \times 21 \text{ h}$)	147 h
Total bandwidth	25 MHz
Central frequency	1406.7 MHz
Spectral resolution	6.104 kHz
Flux/bandpass calibrator	3C138, PKS 1934-638
Phase/gain calibrator	1018-317
Image property	
Mosaic area	$\sim 4.4 \text{ deg}^2$
Velocity coverage	$1159\text{--}4628 \text{ km s}^{-1}$
Velocity resolution	15.5 km s^{-1}
Synthesized beam	$3.7 \text{ arcmin} \times 3.0 \text{ arcmin}$
Final rms	$0.97 \text{ mJy beam}^{-1}$
Velocity resolution	31.0 km s^{-1}
Synthesized beam	$3.3 \text{ arcmin} \times 3.1 \text{ arcmin}$
Final rms	$1.2 \text{ mJy beam}^{-1}$

2.1 Data calibration

Data calibration was done in CASA 4.1.0 (McMullin et al. 2007) using standard methods. Each observing session was flagged and calibrated individually at full spectral resolution. We split off the central 2700 channels, averaged every three channels, and flagged again where necessary. KAT-7 does not Doppler track, so we applied a velocity correction to the barycentric reference frame as described in Carignan et al. (2013). We combined the 19 data sets and averaged a further four channels, such that the spectral resolution of the final data is 15.5 km s^{-1} .

Over the course of the 5 months in which we observed, Antlia went from night time observations, to day time observations. In this period, we noticed an increase in the amount of flagging required, particularly along short baselines, due to the impact of solar interference. However, the fraction of the data flagged was low and although the solar interference was broad-band, it primarily affected data on the lower frequency edge of the band, which was mostly discarded when we split the data.

We performed continuum subtraction with a first-order polynomial to the visibilities using the UVCONTSUB task in CASA. We imaged the cluster using the ‘mosaic’ imager mode in the CLEAN task with robust parameter of 0, and produced two data cubes of 15 and 31 km s^{-1} spectral resolution. The final mosaics cover a roughly circular region 155 pixels in diameter, corresponding to $\sim 1.7 \text{ Mpc}$ at the distance of Antlia, and spanning 228 and 114 channels, respectively. The final restoring beams are 3.7×3.0 and $3.3 \times 3.1 \text{ arcmin}$ corresponding to a spatial resolution of 41×33 and $36 \times 34 \text{ kpc}$, respectively.

2.2 Additional processing: $u = 0$

Despite careful inspection and flagging of the visibilities, after significant averaging, low level artefacts remained in the image cubes. Spatially, we saw horizontal stripes in every channel whose amplitude was comparable to the noise. The dominant mode had a characteristic size of order 1440 arcsec peak-to-peak, corresponding to $\sim 30 \text{ m}$ projected baseline. In the spectral dimension, they appeared as a sinusoidal variation whose phase varied as a function of position in the mosaic. The period of the ripples was of order

¹ <https://science.nrao.edu/facilities/vla/docs/manuals/obsguide/modes/mosaicking>

4 MHz. We believe the artefacts are two manifestations of the same issue, known as the ‘ $u = 0$ problem’.

Interferometer baselines sample varying regions of the UV-plane throughout an observation. When baselines cross the $u = 0$ axis, the fringe rate is equal to zero so incoming signal from the surrounding environment, which is usually added incoherently because of the geometric delay between antennas, is correlated. Anything that is bright enough to be seen in the visibility data is flagged during calibration, so what is left in the image cube is very low level radio frequency interference (RFI) that adds up to be visible after a significant amount of integration time or averaging. This problem affects short baselines more than long baselines, and can be seen in KAT-7 observations after as little as about 8 hours. The source of RFI can be internal or external to the telescope itself, and it is now under investigation how to best mitigate it in MeerKAT while maintaining both interferometric and beam forming mode capabilities.

The solution to remove these artefacts was to Fourier transform (FT) each of the image cubes, mask the ± 2 pixels on either side of the $u = 0$ axis in the UV-plane, and perform the reverse FT back to the image plane. The first FT produces a real and an imaginary UV cube that corresponds to the real and imaginary part of the gridded visibilities. The inverse FT, after masking, produces an amplitude and a phase image cube, but the phases are expected to be zero in the planar approximation because emission from the sky is real. In the resulting image cubes, the horizontal stripes were successfully removed, and the sinusoidal variation of the bandpass was significantly reduced. The rms noise improved by 26 per cent.

2.3 Bandpass fitting, cleaning, and imaging

Due to the detection nature of our observations, we did not know a priori which channels were line-free. Therefore, we performed a second fit to the bandpass in the image plane. The fit was performed by masking pixels in the cubes whose absolute values were greater than 2σ , thereby removing emission and absorption features from influencing the fit. The pixel columns were then fitted with a spline function which was subtracted from the image, removing residual structure in the bandpass.

Finally, bright sources were cleaned in the standard way to remove sidelobes. Overall, the rms noise improved by 29 per cent. In the 15 km s^{-1} cube, we achieved a final value of $0.97 \text{ mJy beam}^{-1} \text{ channel}^{-1}$, corresponding to a 3σ H I mass sensitivity of $4.49 \times 10^7 M_{\odot}$ over 45 km s^{-1} . In the 31 km s^{-1} cube, the rms noise improved by 32 per cent and we achieved a value of $1.2 \text{ mJy beam}^{-1} \text{ channel}^{-1}$, corresponding to $7.4 \times 10^7 M_{\odot}$ over 60 km s^{-1} .

3 ANCILLARY DATA

To compliment the radio observations, we compiled spectroscopic redshifts from the literature and online for all objects in the field, and used the NASA Extragalactic Database (NED)² to cross-match H I detections with optical, infrared, and ultraviolet sources. In addition, we have extracted fluxes from the *WISE* survey (Wright et al. 2010) for extended sources in the KAT-7 mosaic.

² The NASA/IPAC Extragalactic Database (NED) is operated by the Jet Propulsion Laboratory, California Institute of Technology, under contract with the National Aeronautics and Space Administration.

3.1 Optical velocity measurements

The velocities for the brightest optical cluster members come from Ferguson & Sandage (1990) who compiled the work of Lauberts (1982), Hopp & Materne (1985), and Sandage & Tammann (1987). Velocity information on an additional 25 early-type galaxies comes from Smith Castelli et al. (2008) who added velocities from 6dF Galaxy Survey (6dFGS; Jones et al. 2004) and their own spectroscopic measurements. Smith Castelli et al. (2012) observed an additional 22 dwarf elliptical and dwarf spheroidal galaxies from the central region of Antlia. Finally, we have included 28 more unique redshifts from the final 6dFGS catalogue³ (Jones et al. 2009) within a roughly 1.4 radius from NGC 3268 in the redshift range of the cluster, $1200\text{--}4200 \text{ km s}^{-1}$ (Smith Castelli et al. 2008).

In total, we find 97 cluster members with optical redshift measurements in the literature and online. For sources with multiple redshifts measurements, we use those from Smith Castelli et al. (2008) or 6dFGS since they provide the largest uniform sample. This results in a slightly narrower velocity dispersion derived for the cluster in Section 4.2 than if we preferentially use Ferguson & Sandage (1990) redshifts. The estimated uncertainty for individual 6dF redshifts is better than 55 km s^{-1} (Jones et al. 2009).

3.2 WISE data

WISE is a confusion limited survey which observed the entire sky in four mid-infrared bands: 3.4, 4.6, 12, and $22 \mu\text{m}$, conventionally known as *W1*, *W2*, *W3*, and *W4* (Wright et al. 2010). The *W4* band performance has recently been updated (Brown, Jarrett & Cluver 2014) and we apply the new bandwidth performance in which the central wavelength is closer to $23 \mu\text{m}$. In the local (rest) Universe, the *WISE* bands are sensitive to: evolved stars, silicate absorption, polycyclic aromatic hydrocarbons (PAH) molecular bands associated with star formation, warm ISM dust from reprocessed star formation and active galactic nucleus (AGN) disc accretion, respectively (Cluver et al. 2014, and references therein). The $3.4\text{--}4.6 \mu\text{m}$ colour has proven to be a powerful tool to measure stellar mass in the nearby Universe (Jarrett et al. 2013; Cluver et al. 2014), and in particular the $12 \mu\text{m}$ flux is a relatively useful estimate of star formation rate in star-forming galaxies (see also Donoso et al. 2012). It is most reliable at high rates of dust-obscured star formation, but shows correlation to low star formation rates (relative to $\text{H}\alpha$), albeit with increasing scatter.

Large image mosaics ($\sim 3 \text{ deg}^2$) of the KAT-7 field were constructed from *WISE* individual image frames and using the ‘drizzle’ method described in Jarrett et al. (2012). The resulting images are more sensitive to resolved-galaxy detection and characterization compared to the resolution degraded public-release *WISE* mosaics. Given the close proximity of the Antlia Galaxy Cluster and the ~ 6 arcsec angular resolution of the *W1* mosaics, we expect most cluster members to be resolved by *WISE*.

The *WISE* detection sensitivity depends on the number of orbit crossings. In the case of the Antlia field, the *W1* and *W2* bands have a coverage of ~ 27 orbits, and the longer bands have about half of that amount. Accordingly, in the 3 deg^2 area *WISE* detects approximately 480 extended sources in the most sensitive *W1* band using the pipeline developed for *WISE* galaxy characterization (Jarrett et al. 2013). In the area covered by the KAT-7 mosaic, *WISE* detects approximately 240 resolved sources. These include both Antlia and background galaxies. The *W1* detections are limited to a

³ <http://www-wfau.roe.ac.uk/6dFGS/>

signal-to-noise ratio, $S/N = 10$, reaching a limiting Vega magnitude of 16.05 (117 μJy). The other bands, carried by the *W1* detection, have extractions that are limited to $S/N = 3$, and the corresponding limiting magnitudes are respectively (*W2*, *W3*, *W4*): 15.8 (82 μJy), 12.9 (200 μJy) and 8.9 (2.3 mJy). At the bright end, the largest and brightest galaxy detected is the early-type NGC 3268, with a *W1* 1σ isophotal diameter of 6.83 arcmin and *W1* integrated flux of 7.76 mag (241 mJy). We use this resolved-galaxy catalogue to position cross-match with the KAT-7 H I detections, and the *AllWISE* point source catalogue for faint Antlia galaxies not resolved by *WISE*.

4 RESULTS

We report 37 detections in H I with KAT-7 in the Antlia Cluster. Of these 35 are new H I detections, and 27 are new redshift detections at any wavelength. Among the H I sources, 24 are strongly detected in H I emission in both the 15 and 31 km s^{-1} cubes, one is strongly detected in absorption, five are considered more marginal detections because they are relatively low signal-to-noise and appear in only 3–4 channels in the 15 km s^{-1} cube, and seven were only detected in the 31 km s^{-1} cube. The H I detections are completely or mostly unresolved. From comparison with Digitized Sky Survey (DSS) optical images, most of our sources do not appear to suffer from blending or confusion. Only one detection (KAT7HI J103009-360601) has two good optical galaxies within the KAT-7 beam, neither of which have an alternate redshift measurement, but both of which were identified as potential cluster members by Ferguson & Sandage (1990). We list the brighter of the two optical galaxies as the most likely optical counterpart.

Candidate detections were identified by examining cubes channel by channel, and then confirming their profile in position–velocity slices. Detections were required to be at least 3σ and spatially contiguous at that level in at least three adjacent velocity channels in the 15 km s^{-1} cube and two adjacent velocity channels in the 31 km s^{-1} cube. Additionally, we searched optical (DSS), ultraviolet (*GALEX*), and infrared (*WISE*) images for each candidate H I detection to determine the most likely stellar counterparts. We rejected marginal candidates if they both lacked a stellar counterpart, and had low H I signal in noisy channels. Two objects (KAT7HI J103100-350027, KAT7HI J102657-342309) lack a stellar counterpart in NED or *WISE*, but the first has a strong, well-defined H I profile in clean channels (see the Appendix), and the second is spatially convincing albeit with a weaker profile. KAT7HI J103222-345033 lacks a counterpart entry in NED or *WISE*, but has a faint apparently star-forming optical counterpart in DSS images. Unfortunately, it is confused by a bright star diffraction spike in *WISE W1* images.

Table 2 describes the H I properties of the detections and their best known stellar counterpart. Column (1) is our catalogue number, and superscripts indicate a note on the H I detection. Columns (2) and (3) are the H I centroid in J2000 coordinates. The centroid was determined by the peak of a two-dimensional Gaussian fit to the total intensity map of the source. Column (4) is the intensity weighted velocity of the galaxy at the H I centroid. The uncertainty on these values is roughly 22 km s^{-1} . Column (5) is the total integrated flux. Column (6) is the rms in the H I total intensity (moment 0) map. Column (7) is the H I mass, assuming all galaxies lie at the distance of the cluster, $D = 38.1 \text{ Mpc}$. Column (8) is the best optical counterpart, determined by manual inspection of DSS and *GALEX* images. Column (9) is the optical offset between the H I centroid and the optical counterpart in arcmin. Column (10) is

the candidate cluster member from the Ferguson & Sandage (1990) Antlia catalogue. Column (11) is the velocity measured from optical spectroscopy.

4.1 Spatial distribution of H I detected objects

The overall H I distribution of galaxies reflects the effect of the cluster environment (Bravo-Alfaro et al. 2000), and reveals the most recent accretion of galaxies and groups by the cluster. Fig. 1 shows contours of the 37 H I detections from the KAT-7 mosaic overlaid on a *WISE W1* image centred on NGC 3268. To give an impression of the three-dimensional structure, the contours are colour coded by the systemic velocity of the galaxy.

Several points are immediately evident from Fig. 1. First, H I detections avoid the immediate volume around NGC 3268. The closest H I detections appear in a ring surrounding the elliptical galaxy at projected distances greater than 200 kpc. This is in contrast to the elliptical galaxy NGC 3258 to the south-east which has four H I detected galaxies projected within 80 kpc. Secondly, we observe that galaxies which are blueshifted with respect to the cluster velocity preferentially lie in the ring, whereas redshifted galaxies are distributed asymmetrically and preferentially to the north. In the context of the large-scale structure, this is in the same direction as the Hydra Cluster which, at a distance of 55 Mpc ($v = 3777 \text{ km s}^{-1}$) and a projected $\sim 8^\circ$ from Antlia, lies $\sim 18 \text{ Mpc}$ away. In the following sections, we argue that the distribution of H I is due to asymmetry in the accretion from the surrounding environment. It may be suggestive of accretion along a filament connecting the two clusters.

Fig. 2 shows that the ring of H I detections is coincident with the extent of the X-ray halo around NGC 3268. The X-ray intensity shows complex structure. On the large scale, it is elongated in the direction between NGC 3268 and NGC 3258, which has its own extended X-ray emission. The H I detections avoid the brightest X-ray regions. We conclude that NGC 3268 and its large X-ray halo are the dominant and central structure of the Antlia Cluster, while NGC 3258 and its globular cluster and dwarf galaxy population are likely to be more recently arrived. Consistent with observations of other nearby clusters (Virgo, Chung et al. 2009; Coma, Bravo-Alfaro et al. 2000), the centre of Antlia is very H I deficient, and suggests the extended hot X-ray halo is, in part, responsible for removing H I from galaxies.

4.2 Velocity distribution and infall

Fig. 3 shows a velocity histogram of cluster members with only optical redshifts (predominantly early-type galaxies), versus those with H I redshifts (predominantly late-type galaxies; Gavazzi et al. 2008). We find that the H I detected galaxies are by no means well fit by a Gaussian and may be better described as multimodal (Scodreggio et al. 1995). None the less, as a population they have a broader velocity distribution than the optical sample, by a factor of 2.31, and are centred at a systematically higher redshift by 700 km s^{-1} .

Assuming isotropic orbits, the ratio between the velocity dispersion of an infalling galaxy population and a virialized population is predicted to be $\sigma_{\text{infall}}/\sigma_{\text{vir}} = \sqrt{2} = 1.4$ (Conselice, Gallagher & Wyse 2001). For example, the ratio of velocity dispersion between Coma late-type and early-type galaxies is 1.4, while the ratio between Virgo late- and early-type galaxies is 1.64 (Table 3). Virgo also shows greater asymmetry in the overall distribution of galaxies and in its X-ray contours than Coma, suggesting it is

Table 2. Antlia Cluster H I Source Catalog.

HI catalogue number	RA (J2000)	Dec. (J2000)	Velocity (km s ⁻¹)	S_{int} (Jy km s ⁻¹)	σ (Jy beam ⁻¹ km s ⁻¹)	log $M_{\text{H I}}$ (M_{\odot})*	Optical Counterpart†	Opt offset (arcmin)	Ferguson & Sandage (1990)	Opt vel (km s ⁻¹)
1	10:30:53	-34:53:44	4108	2.12	0.117	8.86	J103054.98-345324.1	0.525	FS90 248	
2	10:31:48	-35:12:07	4037	0.89	0.094	8.48	LEDA 083201	0.356	FS90 297	
3	10:29:15	-35:41:39	3903	0.35	0.055	8.08	J102915.11-354141.6	0.071	FS90 137	
4	10:26:25	-34:57:58	3358	1.53	0.111	8.72	NGC 3249	0.590	FS90 028	3386 ³
5	10:27:48	-34:32:10	3214 ¹	5.22	0.100	9.25	ESO 375- G 030	0.336		
6	10:31:32	-35:32:06	3228	0.51	0.073	8.24	J103133.77-353217.1	0.366	FS90 284	
7	10:28:03	-35:27:00	3195	0.30	0.056	8.00	2MASX J10280313-3526317	0.472	FS90 068	3188 ³
8	10:31:00	-35:00:27	3142	1.49	0.092	8.71				
9	10:30:00	-34:47:27	2996	0.42	0.051	8.16	J102957.52-344831.3	1.187		
10	10:26:13	-35:08:26	2952	0.31	0.043	8.03	J102612.07-350852.5	1.014		
11	10:28:32	-35:27:34	2895	0.55	0.076	8.27	LEDA 083019	0.416	FS90 098	2890 ⁴
12	10:31:23	-35:13:10	2549	1.00	0.106	8.53	NGC 3258C	0.243	FS90 277	2597 ⁵
13 ^b	10:31:29	-35:18:26	2582	0.24	0.059	7.92	J103134.49-351920.4	0.526		
14 ^b	10:31:07	-35:31:08	2648	0.17	0.043	7.77				
15	10:28:31	-35:40:41	2450	0.99	0.063	8.52	J102831.57-354036.9	0.199	FS90 093	
16	10:30:19	-35:35:37	2365	1.01	0.086	8.54	J103020.84-353533.2	0.336	FS90 212	
17 ^b	10:28:42	-35:12:38	2458	0.22	0.039	7.87	J102840.97-351234.9	0.251		
18	10:28:44	-35:06:49	2419	0.40	0.071	8.14	LEDA 083023	0.609	FS90 102	
19	10:29:26	-35:00:57	1829	0.87	0.065	8.47	J102926.43-350052.6	0.138	FS90 147	
20	10:29:10	-35:41:20	2147	1.04	0.065	8.55	MCG -06-23-034	0.225	FS90 131	2104 ⁴
21	10:29:02	-35:34:28	2620	0.32	0.073	8.04	FS90 120	0.392	FS90 120	
22	10:29:24	-34:40:50	4117	0.75	0.110	8.41	2MASX J10292838-3440214	1.020	FS90 152	4093 ³
23	10:28:43	-34:41:54	4267	0.58	0.096	8.30	LEDA 083022	0.182	FS90 101	
24 ^a	10:31:51	-34:51:17	3359 ²	-1.59	0.135		NGC 3281	0.232	FS90 300	3437 ³
25	10:27:00	-36:13:48	3079	3.57	0.146	9.09	ESO 375- G 026	0.514	FS90 029	3087 ³
26 ^b	10:30:13	-34:29:28	4381	0.536	0.066	8.26	J103013.38-343103.4	1.593		
27	10:32:22	-34:50:33	4287	1.25	0.118	8.63				
28	10:30:19	-34:24:31	2990	0.82	0.092	8.45	ESO 375- G 047	0.249		3068 ³
29 ^b	10:30:54	-35:58:25	3235	0.30	0.037	8.00	J103050.43-355836.5	0.747		
30	10:28:37	-36:26:00	3438	1.17	0.026	8.60	LEDA 083018	0.562	FS90 097	
31	10:32:55	-36:04:12	4145	0.47	0.072	8.21	LEDA 083229	0.501	FS90 330	
32 ^b	10:32:13	-34:32:10	3315	2.89	0.162	9.00	J103213.22-343158.2	0.201		
33 ^b	10:31:00	-34:23:51	3794	3.02	0.147	9.01	J103102.44-342410.1	0.596		
34	10:30:09	-36:06:01	2241	0.67	0.123	8.36	J103006.07-360624.3	0.723	FS90 197	
35	10:28:01	-35:18:54	4124	0.45	0.147	8.19	LEDA 082990	0.273	FS90 065	
36	10:26:55	-34:39:07	3904	4.38	0.258	9.18	LEDA 082948	0.914	FS90 021	
37	10:26:57	-34:23:09	3797	3.91	0.166	9.13				

*Assumes all galaxies are at the same cluster distance, 38.1 Mpc. †Optical counterparts that only consist of a J2000 coordinate come from the GALEXASC catalogue. ^aDetected in absorption. ^bConsidered a more marginal H I detection. ¹Detected by HIPASS at 3242 km s⁻¹. ²Reported by Nançay in emission at 3200 km s⁻¹ (see text). ³6dF (Jones et al. 2009). ⁴Smith Castelli et al. (2008). ⁵Ferguson & Sandage (1990).

further from being virialized. Similar ratios have been seen in larger samples of galaxy clusters when characterizing galaxies by colour or morphology (e.g. Adami, Biviano & Mazure 1998, and references therein). The value of Antlia is much larger than Virgo or Coma and is consistent with the H I galaxies being a younger cluster population with regards to their relative age since accretion than those galaxies which are not detected in H I.

The offset in the H I velocity distribution from the optical population is likely due to a real asymmetry in the accretion of galaxies from the surrounding environment. As an example, Gavazzi et al. (1999) used the Tully–Fisher relation to show that redshifted galaxies in Virgo are closer than blueshifted galaxies: galaxies on the far side of the cluster are falling inward – towards us – and thus are blueshifted with respect to the cluster centre. If the change in observed mass sensitivity between the front and back of the Antlia Cluster were *significant*, we would expect a paucity of galaxies at blueshifted velocities. However, generously assuming Antlia is 2 Mpc in diameter translates to only a 10 per cent difference in H I mass sensitivity. Thus, we argue that the accretion on to Antlia is intrinsically asymmetric, rather than observational bias. Further,

given the larger scale structure around Antlia, we suggest that the asymmetry is due to a higher rate of accretion of galaxies from behind the cluster, notably towards Hydra, than from in front of it.

4.3 Infrared counterparts and properties

Using a radius of ~ 1 arcmin, positional cross-matching between the Antlia H I detections and the *WISE* resolved extractions gives rise to 23 confident matches. Of these, eight are below the standard $S/N = 10$ limit for the *WI* extended source catalogue, but are measured a posteriori based on the H I detection. Another four H I detections were matched with unresolved *WISE* sources from the *AllWISE* catalogue (Cutri et al. 2013), although the chance for a false match is much greater since these *WISE* sources may be foreground stars or background galaxies. Hence, the resolved-galaxy match rate is 23/37 (62 per cent) and the total match rate is 27/37 (73 per cent).

We use *WISE* extracted photometry to estimate the host stellar mass and the obscured star formation rate. For the stellar mass, we

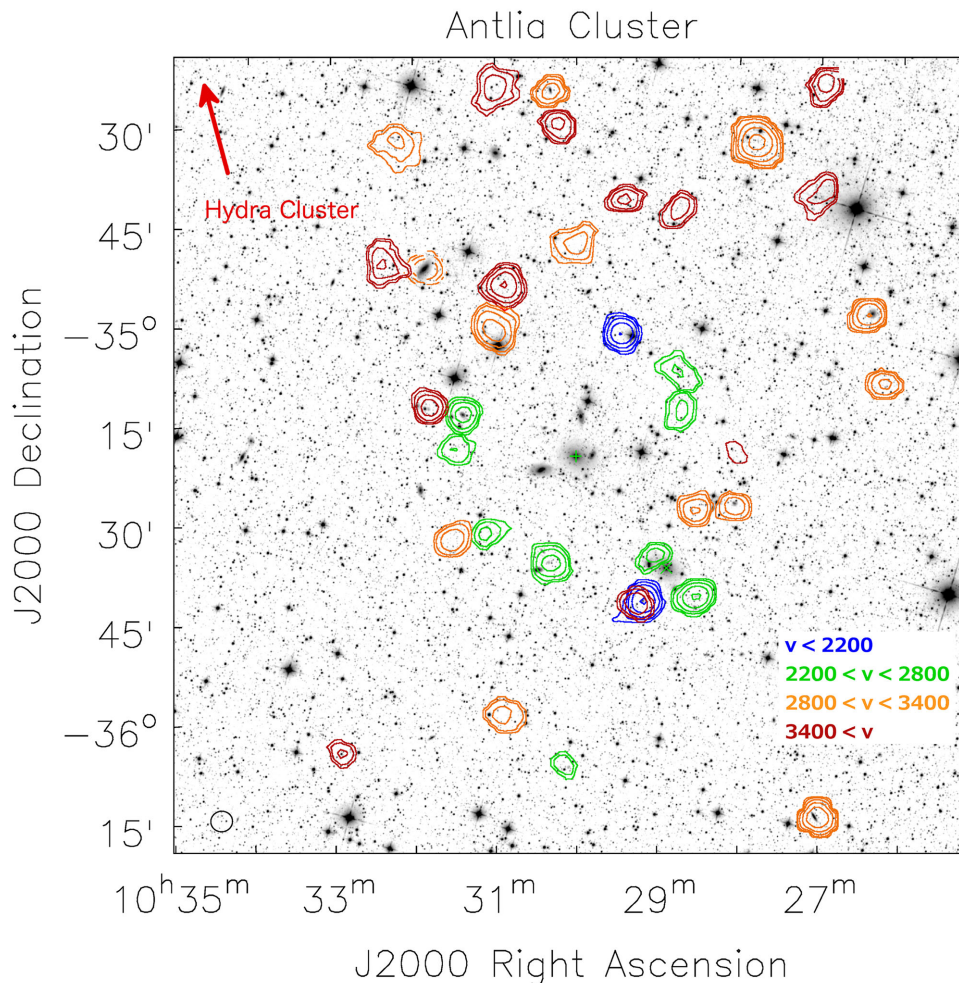


Figure 1. $H\text{I}$ contours on a *WISE* 3.4 μm image. The contours correspond to column densities of 2, 4, 8, 16, 32, $64 \times \sigma$ listed in Column (6) of Table 2 and are colour coded by the systemic velocity of the galaxy. The centre of the image coincides with the centre of the $H\text{I}$ mosaic and the centre of the cluster, NGC 3268. Dashed contours represent the negative $H\text{I}$ absorption around NGC 3281, north-east of the cluster centre.

use the ‘resolved-galaxy’ (i.e. nearby galaxies) M/L relations of Cluver et al. (2014), which in combination with the *W1* 3.4 μm ‘in-band’ luminosities (Jarrett et al. 2013) and the $W1 - W2$ colour (which roughly accounts for metallicity and morphology differences), infers a stellar mass and formal uncertainty through error propagation. Accordingly, for the Antlia matches the minimum stellar mass is found to be $\sim 10^{7.8} M_{\odot}$ (WXSC J10280155-351900.4; or considering point sources, $\sim 10^{6.6} M_{\odot}$ for WISEA J102956.70-344834.1) and the most massive source at $10^{11.1} M_{\odot}$ is the barred lenticular galaxy, NGC 3271.

The star formation rate is estimated from both the 12 and 22 μm measurements using the Cluver et al. (2014) relations. The most sensitive mid-infrared band of *WISE* is the *W3* band, whose rest bandpass captures the 11.3 μm PAH band, a star formation tracer, although exhibiting more scatter in comparison to other indicators (Cluver et al. 2014). Comparatively, *WISE* 22 μm band is less sensitive, but is dominated by dust-emitting continuum in thermal dynamic equilibrium and thus a more reliable SF tracer (e.g. Jarrett et al. 2013). Most of the Antlia sources have very low star formation activity – as traced by the obscured star formation – with most inferred SFR values $< 0.01 M_{\odot} \text{ yr}^{-1}$, consistent with galaxy cluster members in the local Universe. At the high end, a few sources do have rates greater than unity, notably NGC 3281 with an inferred value of ~ 10 and $10.7 M_{\odot} \text{ yr}^{-1}$ in *W3* and *W4*,

respectively. It should be noted that since this host galaxy has an AGN (type II Seyfert) and thus mid-infrared emission arising from the accretion disc (see Fig. 9 for a visual demonstration of powerful nuclear emission), the host galaxy star formation luminosity is overestimated and should be treated with caution. In total, 6/37 (16 per cent) of the $H\text{I}$ detections have star formation rates greater than $0.05 M_{\odot} \text{ yr}^{-1}$.

Table 4 lists the properties of *WISE* counterparts to $H\text{I}$ detected sources: Column (1) is the $H\text{I}$ catalogue number from Table 2; Column (2) is the *WISE* detection; Columns (3)–(6) are the *W1*, *W2*, *W3*, and *W4* Vega magnitudes; Column (7) is the stellar mass; Columns (8) and (9) are the star formation rate calculated from the 12 and 22 μm bands; Column (10) is the $H\text{I}$ mass-to-stellar mass ratio.

4.4 Stellar masses, star formation rates, and gas fractions

Combining quantitative measures of the gaseous and stellar components of cluster members provides the most complete context for studying the overall cluster dynamics and galaxy evolution with environment. From cross-matching to infrared sources and Table 4, we find that the $H\text{I}$ population is dominated (28/37) by dwarf galaxies with $M_{*} < 10^9 M_{\odot}$, and star formation rates below $0.1 M_{\odot} \text{ yr}^{-1}$.

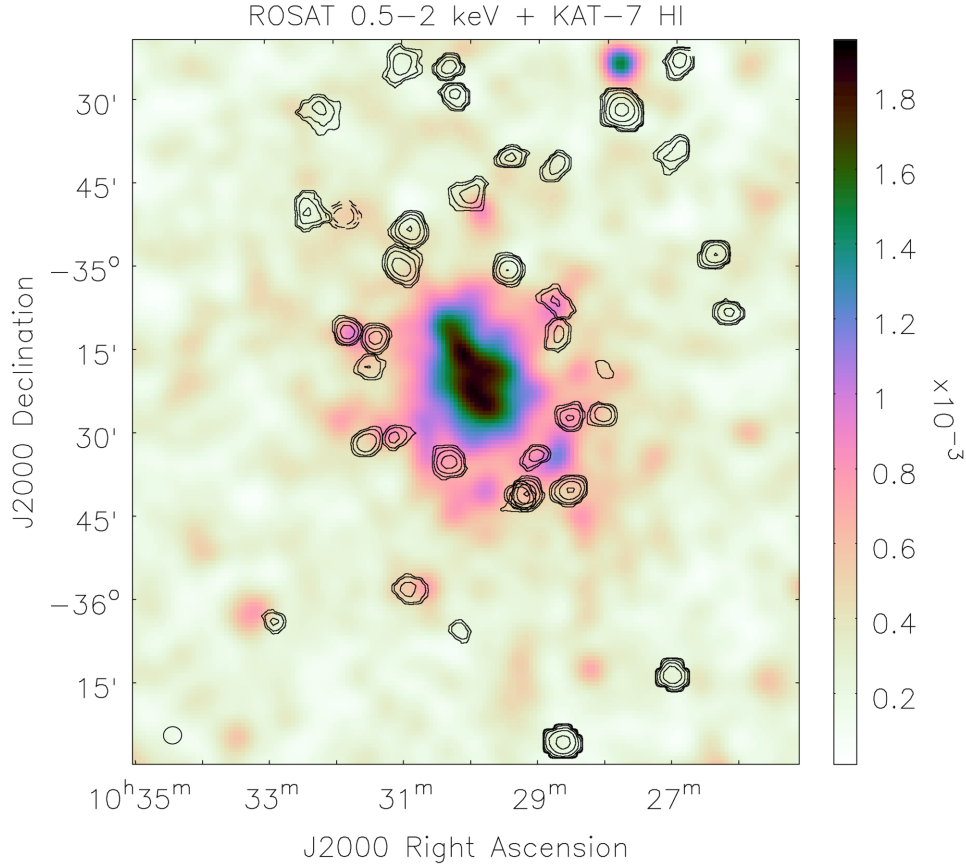


Figure 2. *ROSAT* 0.5–2 keV hard X-ray intensity image from SkyView (<http://skyview.gsfc.nasa.gov/current/cgi/titlepage.pl>) (McGlynn, Scollick & White 1996), convolved with a 5×5 arcmin Gaussian, and overlaid with H I contours. The X-ray image has had no point source removal performed.

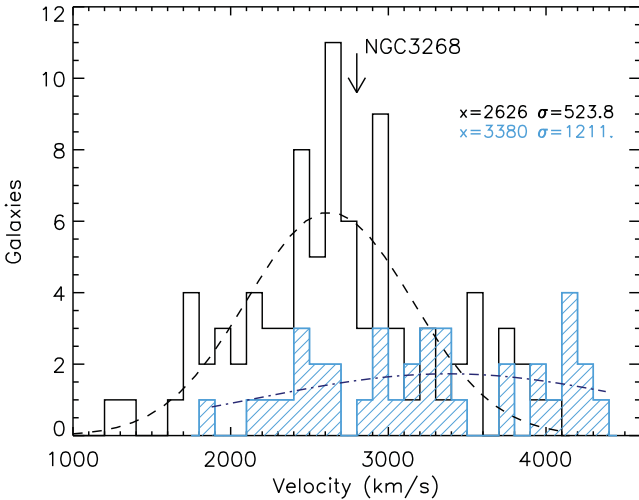


Figure 3. Velocity histograms of 97 galaxies with only optical spectroscopic redshifts (black) and 37 H I spectroscopic redshifts (blue). The optical sample is dominated by early-type galaxies, while the H I sample is primarily late-type galaxies. The legend gives the centre and dispersion of each velocity distribution. The velocities of late-type galaxies are systematically offset and more broadly distributed than early-type galaxies.

Further, most objects have $M_{\text{H I}}/M_* < 1$ and the lowest stellar mass objects have the highest gas fractions, consistent with what is seen on average across all environments (Huang et al. 2012; Maddox et al. 2015).

Table 3. Summary of velocities for Antlia and other clusters.

Description	v (km s $^{-1}$)	σ (km s $^{-1}$)
NGC 3268	2800	
NGC 3258	2792	
Antlia opt galaxies	2626	524
Antlia H I galaxies	3380	1211
All Antlia galaxies	2747	656
Virgo early-type galaxies	1134	573
Virgo late-type galaxies	1062	888
All Virgo galaxies	1112	757
Coma early-type galaxies	6807	1017
Coma late-type galaxies	7375	1408
All Coma galaxies	6917	1038

Notes. Virgo (Binggeli, Tammann & Sandage 1987); Coma (Colless & Dunn 1996).

Fig. 4 shows the spatial distribution of all *WISE* extended sources in and around the Antlia Cluster, including background galaxies. (There are no known foreground galaxies below $v = 1200$ km s $^{-1}$.) The figure conveys a sense of the overdensity of galaxies within Antlia in relation to the H I detections, and we highlight ‘star-forming’ galaxies, relative to the majority of passive cluster members, if they have $W2 - W3 > 1.5$ and $\text{SFR}_{12\mu\text{m}} > 0.05 M_{\odot} \text{ yr}^{-1}$.

We find 13 of the 124 spectroscopically confirmed cluster members (10 per cent) and 6/37 H I detections (16 per cent) have star formation rates greater than $0.05 M_{\odot} \text{ yr}^{-1}$. At least seven cluster members have appreciable star formation without detected

Table 4. *WISE* counterparts to H I detections in the Antlia Cluster.

H I catalogue number	WISE name (J2000)	3.6 μ m (W1 mag)	4.5 μ m (W2 mag)	12 μ m (W3 mag)	22 μ m (W4 mag)	Stellar mass $\log(M_{\odot})$	12 μ m SFR ($M_{\odot} \text{ yr}^{-1}$)	22 μ m SFR ($M_{\odot} \text{ yr}^{-1}$)	H I mass fraction $M_{\text{H I}}/M_{*}$
1†	WXSC J10305496-345325.0	13.94 \pm 0.03	13.85 \pm 0.07	13.44 \pm 0.49		8.48 \pm 0.17	<0.01		2.4
2†	WXSC J10314945-351220.0	13.35 \pm 0.02	13.30 \pm 0.06	12.80 \pm 0.33	10.14 \pm 0.74	8.80 \pm 0.13	<0.01	0.01	0.48
4	WXSC J102622.21-345748.8	9.81 \pm 0.01	9.74 \pm 0.02	6.49 \pm 0.02	4.83 \pm 0.04	10.19 \pm 0.06	0.98 \pm 0.02	0.73 \pm 0.02	0.03
5†	WXSC J10274705-343156.0	14.07 \pm 0.04	13.97 \pm 0.09	15.22 \pm 1.97	11.40 \pm 1.99	8.41 \pm 0.20	<0.01	0.01	6.9
7	WXSC J102803.15-352631.4	11.50 \pm 0.01	11.47 \pm 0.03	8.35 \pm 0.03	6.81 \pm 0.11	9.61 \pm 0.07	0.15 \pm 0.01	0.17 \pm 0.01	0.02
10†	WXSC J10261230-350856.0	13.66 \pm 0.03	13.64 \pm 0.06	12.14 \pm 0.22		8.77 \pm 0.13	<0.01		0.18
11	WXSC J102834.19-352740.0	13.51 \pm 0.02	13.61 \pm 0.05	11.49 \pm 0.23	> 8.37	9.12 \pm 0.12	0.01		0.14
12	WXSC J103124.21-351313.8	10.50 \pm 0.01	10.47 \pm 0.02	7.43 \pm 0.02	5.92 \pm 0.07	10.01 \pm 0.06	0.39 \pm 0.01	0.32 \pm 0.01	0.03
14*	WISEA J103110.89-353125.3	15.66 \pm 0.04	15.69 \pm 0.12	> 12.33	> 8.88	8.10 \pm 0.24			0.46
15†	WXSC J10283170-354038.0	13.71 \pm 0.03	13.65 \pm 0.07	13.29 \pm 0.42		8.64 \pm 0.16	<0.01		0.76
16	WXSC J103020.84-353530.8	13.32 \pm 0.02	13.22 \pm 0.06	11.52 \pm 0.19	> 8.33	8.73 \pm 0.13	0.01		0.65
18	WXSC J102843.69-350724.9	13.98 \pm 0.03	14.03 \pm 0.07	> 12.20	> 8.84	8.82 \pm 0.16			0.21
19	WXSC J102926.42-350052.5	13.98 \pm 0.03	13.73 \pm 0.06	11.70 \pm 0.22	8.88 \pm 0.42	8.06 \pm 0.15	0.01	0.03	2.6
20	WXSC J102911.05-354115.5	11.85 \pm 0.01	11.72 \pm 0.03	8.25 \pm 0.03	6.55 \pm 0.08	9.21 \pm 0.07	0.17 \pm 0.01	0.20 \pm 0.01	0.22
21†	WXSC J10290171-353358.0	14.60 \pm 0.05	14.85 \pm 0.13	13.53 \pm 0.69		8.83 \pm 0.25	<0.01		0.16
22	WXSC J102928.40-344021.7	12.42 \pm 0.01	12.40 \pm 0.03	10.59 \pm 0.08	> 8.56	9.26 \pm 0.08	0.02		0.14
23*	WISEA J102842.68-344145.5	16.17 \pm 0.06	16.03 \pm 0.15	> 12.62	> 9.22	7.47 \pm 0.29			6.8
24†	WXSC J103152.08-345113.2	7.91 \pm 0.01	7.12 \pm 0.02	4.09 \pm 0.02	1.37 \pm 0.03	9.60 \pm 0.05	10.69 \pm 0.21	10.05 \pm 0.09	0.10
25	WXSC J102702.45-361341.1	9.59 \pm 0.01	9.43 \pm 0.02	5.85 \pm 0.02	4.08 \pm 0.03	10.07 \pm 0.05	1.85 \pm 0.04	1.30 \pm 0.03	6.5
26*	WISEA J103014.11-342844.9	15.13 \pm 0.04	14.82 \pm 0.07	12.52 \pm 0.46	> 8.62	7.45 \pm 0.16	<0.01		0.05
28	WXSC J103019.54-342417.2	10.73 \pm 0.01	10.62 \pm 0.02	7.21 \pm 0.02	5.54 \pm 0.04	9.71 \pm 0.06	0.48 \pm 0.01	0.43 \pm 0.01	0.23
31	WXSC J103254.69-360436.9	14.08 \pm 0.03	14.15 \pm 0.08	12.25 \pm 0.32	> 8.57	8.85 \pm 0.18	<0.01		65
32*	WISEA J10321325-343157.4	15.13 \pm 0.04	14.72 \pm 0.06	12.10 \pm 0.35	> 9.04	7.19 \pm 0.14	<0.01	0.02	1.62
33	WXSC J10310256-342406.0	13.79 \pm 0.02	13.80 \pm 0.06	11.09 \pm 0.12	9.66 \pm 0.60	8.80 \pm 0.13	0.01		0.29
34	WXSC J10300598-360623.0	13.16 \pm 0.02	13.12 \pm 0.05	11.78 \pm 0.18		8.90 \pm 0.12	0.01		2.43
35†	WXSC J10280155-351900.4	16.21 \pm 0.06	16.88 \pm 0.34	> 12.73	> 9.06	7.80 \pm 0.46	<0.01		18
36†	WXSC J10265082-343845.4	15.41 \pm 0.04	15.33 \pm 0.09	11.90 \pm 0.28	> 8.96	7.91 \pm 0.22	0.01	0.03 \pm 0.01	

† W1 detection with S/N < 10. * Detected in the ALLWISE point source catalogue. † Seyfert II galaxy detected in H I absorption (see Figs 7 and 8). Star formation rates less than 0.03 have measured uncertainties less than 0.005.

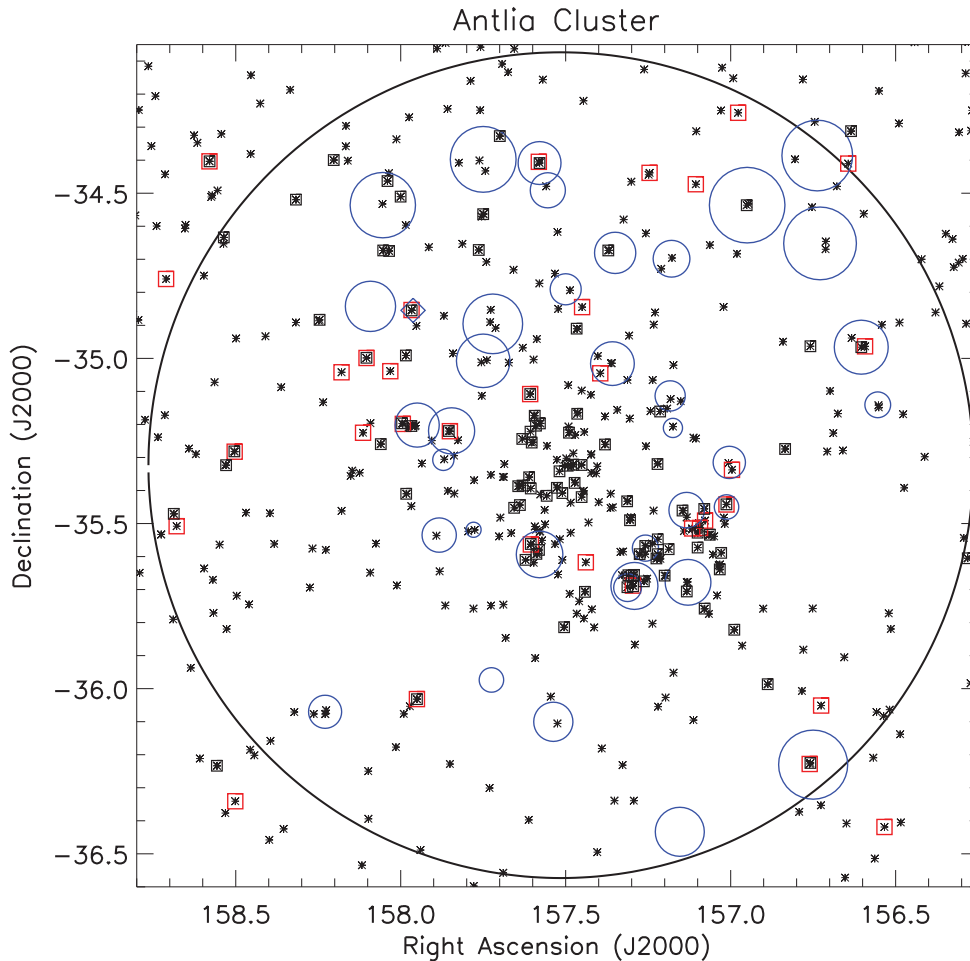


Figure 4. The spatial distribution of *WISE* extended sources in and around the Antlia Cluster. Red symbols highlight ‘star-forming’ over passive galaxies: those with $W2 - W3 > 1.5$ and $SFR > 0.05 M_{\odot} \text{ yr}^{-1}$. Black squares indicate known cluster members with optical redshifts. Blue symbols indicate H I cluster members where the radius is scaled by $\log(M_{\text{H I}})$. The blue diamond is NGC 3281, detected in H I absorption. The blue symbols and black squares make up the subset of galaxies in Fig. 7. The large circle corresponds to the extent of the KAT-7 H I mosaic.

H I. Only one, WXSC J103025.79-350628.7 (MCG -06-23-043; $v = 1781 \text{ km s}^{-1}$, 6dF), is projected within the 200 kpc H I ring: forming stars at a rate of $0.18 M_{\odot} \text{ yr}^{-1}$ in a $10^{9.8} M_{\odot}$ galaxy. Fig. 5 shows the specific star formation rate versus stellar mass for all *WISE* extended sources: most massive galaxies are passive, non-star-forming systems. There are no obvious trends between H I mass, stellar mass, or star formation rate.

Finally, Fig. 6 shows the H I mass versus stellar mass of the Antlia galaxies compared to the mean values of the entire Arcibo Legacy Fast ALFA (ALFALFA) survey (Haynes et al. 2011) sample from Maddox et al. (2015). The ALFALFA sample is dominated by field galaxies (Hess & Wilcots 2013), so this provides a baseline for the impact of the cluster environment on the gas content of individual galaxies compared to the field. Compared to galaxies of the same stellar mass, all Antlia galaxies, except a handful of the dwarfs with $M_{*} < 10^{8.5} M_{\odot}$, are H I deficient and fall well below the 1σ distribution. The most massive galaxies by stellar mass are also the most gas deficient compared to the field, by 2–3 orders of magnitude in H I mass. Meanwhile, the most H I rich, both by mass and gas fraction, are dwarf galaxies. From Fig. 4, we see that these dwarfs reside on the extreme outskirts of the cluster. Their reservoir of H I gas will not survive once they have fallen into the cluster environment.

We conclude that gas stripping is occurring in the cluster, and is happening at appreciable levels outside the X-ray-emitting halo: out to at least ~ 600 kpc in projected radius from the cluster centre. The H I mass is also a measure of future star formation potential. So in addition to removing H I the cluster environment must impact future star formation, but it is unclear whether accretion on to the cluster enhances star formation before quenching it, and whether quenching occurs slowly (Kauffmann et al. 2004) or rapidly (Rines et al. 2005). Galaxy transformation may begin at several virial radii from the cluster centre, but star-forming galaxies can still exist within the cluster X-ray halo (Caldwell et al. 1993). We show that the tracers of star formation are still visible even when the atomic gas has been exhausted or stripped below detectable limits, with star-forming galaxies potentially drawing from a reserve of dense molecular gas (Kenney & Young 1989; Casoli et al. 1991).

5 DISCUSSION

From the spatial distribution of H I detections we have established that, despite the presence of two massive early-type galaxies at its centre, Antlia’s oldest and most dominant structure is centred on NGC 3268. The velocities of the NGC 3268 and NGC 3258 at 2800 ± 21 and $2792 \pm 28 \text{ km s}^{-1}$, respectively, give no hints as

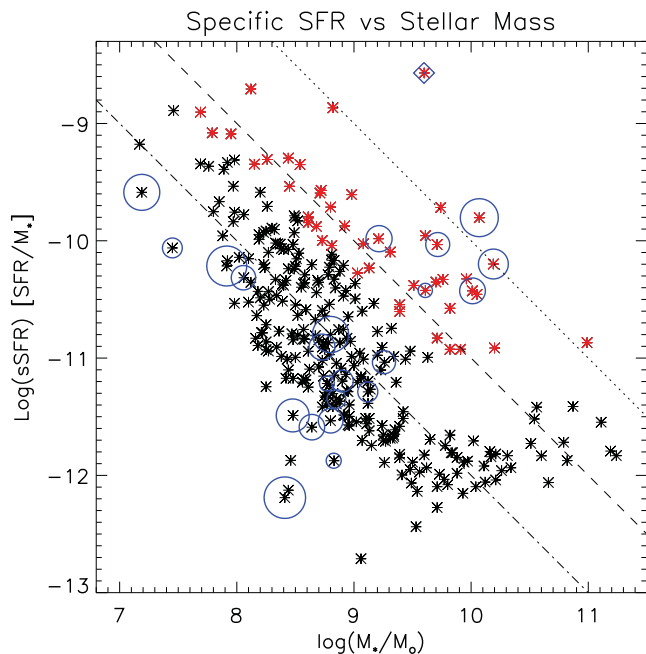


Figure 5. The specific star formation versus stellar mass for all *WISE* extended sources, including Antlia and background galaxies. Star-forming galaxies are in red, and H I detected objects are in blue as in the previous figure. Lines of constant star formation (0.01 , 0.1 , $1 \text{ M}_\odot \text{ yr}^{-1}$) are shown by dot-dashed, dashed, and dotted lines, respectively.

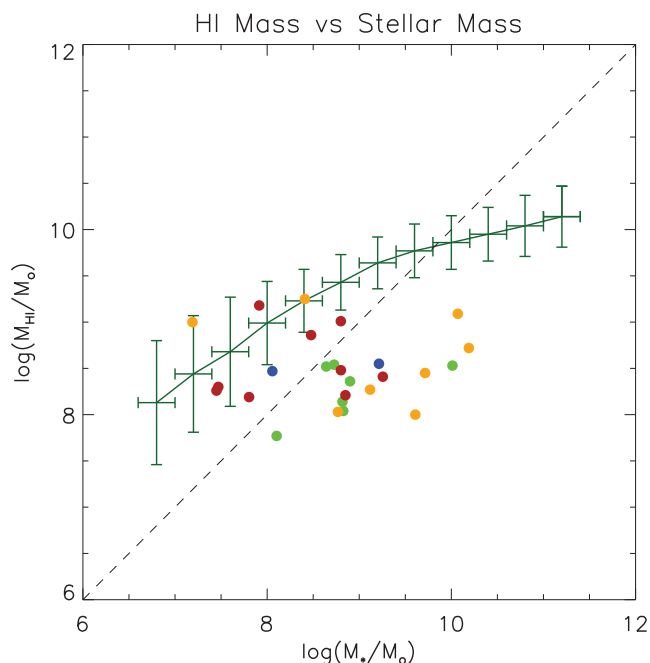


Figure 6. The H I mass versus stellar mass for all H I Antlia galaxies with *WISE* counterparts. The dashed black line is the one-to-one gas fraction relation. The green line is the median values for the ALFALFA survey (Haynes et al. 2011) from Maddox et al. (2015) where the horizontal error bars represent bin size, and the vertical errors bars represent the 1σ distribution. The filled circles are the Antlia H I detections where the colours are the same as Fig. 1.

to how a merger between the two subclusters may have proceeded. Further, we have shown that there is little active star formation going on in the centre of the cluster, and that the H I galaxies within at least 600 kpc of the centre have undergone significant stripping even though they lie outside the X-ray halo.

Statistical dynamical tests in combination with the H I and star formation content of galaxies will help us to further disentangle the assembly history of the cluster. In the following subsections, we analyse the three-dimensional structure of Antlia, investigate substructure in the context of their stellar and neutral gas content, calculate the total cluster mass, and put Antlia into an evolutionary context with respect to other nearby, well-studied clusters. Finally, we present the serendipitous detection of H I absorption in the Compton thick Seyfert II galaxy, NGC 3281.

5.1 Substructure in the Antlia Cluster

The clumpy distribution of galaxies in space and velocity suggests the ongoing accretion and merger of groups with the cluster, in addition to the infall of individual galaxies. The presence of H I or star formation provides a rough timeline for the merger activity because these aspects of infalling galaxies are transformed by the cluster environment. Hopp & Materné (1985) identified five nearby groups to Antlia, which lie at projected distances greater than 2.5 and 1.75 Mpc. We focus on the substructure within the KAT-7 H I mosaic which covers the volume containing 96 per cent of the optical galaxies associated with the cluster (Fig. 4) and extends to 4–5 times the X-ray extent of NGC 3268 subcluster.

Spatial separation is a intuitive quality of substructure, but the DS test provides a measure of the relative kinematic deviation of substructure from the global kinematics (Dressler & Shectman 1988). The test and its application to both groups and clusters is well summarized in Hou et al. (2012). In particular, this test has been used to understand whether infall enhances or quenches star formation and gas accretion on to systems (e.g. Pranger et al. 2013). Recent results find that substructure in groups and clusters is correlated with an increase in active galaxies (Hou et al. 2012; Cohen et al. 2014).

For each galaxy, i , we calculate the mean velocity, \bar{v}_{local}^i , and velocity dispersion, σ_{local}^i , for it and its $N_{\text{nn}} = 10$ nearest neighbours compared to the mean velocity, \bar{v} , and velocity dispersion, σ , of the entire cluster:

$$\delta_i = \left(\frac{N_{\text{nn}} + 1}{\sigma^2} \right) [(\bar{v}_{\text{local}}^i - \bar{v})^2 + (\sigma_{\text{local}}^i - \sigma)^2] \quad (1)$$

$$\Delta = \sum_i \delta_i / n_{\text{members}}, \quad (2)$$

where the threshold $\Delta = 1$ indicates the distribution is close to Gaussian and the fluctuations are random. For Antlia, $\bar{v} = 2747 \text{ km s}^{-1}$ and $\sigma = 656 \text{ km s}^{-1}$ for 124 galaxies with H I or optical redshifts, and we find $\Delta = 1.4$, suggesting the presence of substructure.

Cohesive substructure is identified with the DS statistic by locating associations of galaxies which deviate from the overall cluster systematics in the same way, and are close to each other in spatial proximity and velocity. In Fig. 7, we plot the kinematic deviation, δ_i , for each galaxy as a function of position, and substructures can be identified by collections of circles of the same size and colour. We can then compare these substructures with Fig. 4 to learn if they are gas rich and/or star forming. In the following sections, we

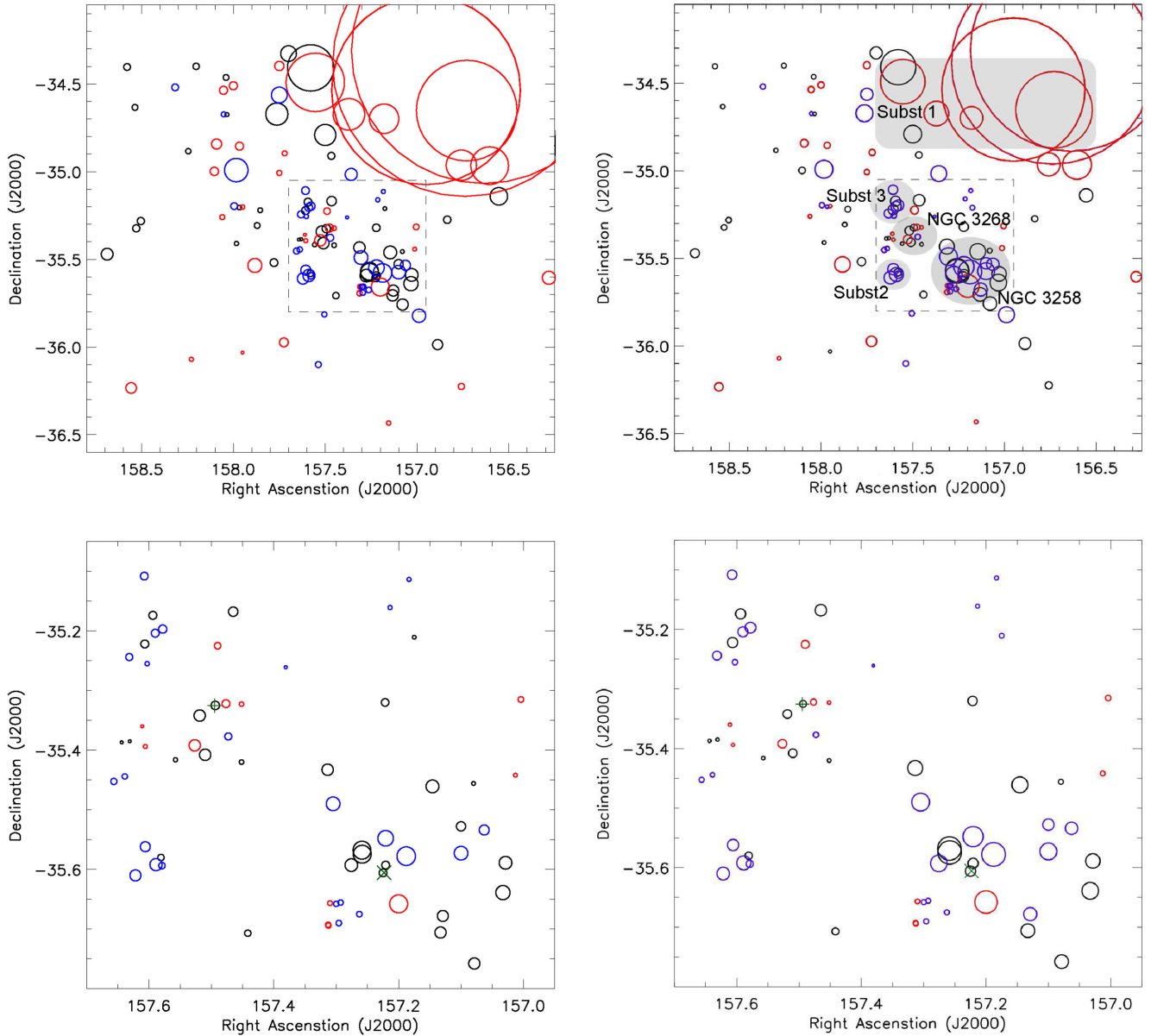


Figure 7. Dressler & Shectman (1988) ‘bubble plot’ for 124 galaxies with redshifts, where the centre of each circle corresponds to the position of a galaxy, and the size scales with $\exp(\delta_i)$. Left: the kinematic deviation has been calculated from the derived mean velocity of the cluster $\bar{v} = 2747 \text{ km s}^{-1}$. Right: the kinematic deviation is calculated assuming $\bar{v} = 2800 \text{ km s}^{-1}$, the velocity of NGC 3268. Black symbols are galaxies with velocities within $\pm 0.5\sigma$ of \bar{v} ; blue symbols are blue shifted to $v < -0.5\sigma$; red symbols are redshifted to $v > +0.5\sigma$. Top plots are the entire cluster; bottom plots are zoomed in on the inset, and show how the choice in \bar{v} brings out substructure within the cluster centre. The ‘+’ denotes the position of NGC 3268 and the nominal cluster centre; the ‘X’ denotes the position of NGC 3258. Substructure is labelled in grey in the upper right panel for clarity.

discuss these plots in detail, considering first the inner virialized region, then the outskirts where infall is reflected in the kinematics.

5.1.1 The inner cluster

The systemic velocities of NGC 3268 and NGC 3258 are 2800 km s^{-1} and 2792 km s^{-1} , respectively, and their similar velocities tell us nothing about the galaxies’ relative roles within the cluster. With the DS test, we identify differences between their respective subclusters. Fig. 7 shows DS test performed twice: once as prescribed above with $\bar{v} = 2747 \text{ km s}^{-1}$ (left-hand panels), and once using the velocity of

NGC 3268 as $\bar{v} = 2800 \text{ km s}^{-1}$ (right-hand panels). Quantitatively, Δ is the same for both choices of \bar{v} , but the change in δ_i for individual galaxies exposes local kinematic variation between the NGC 3268 and NGC 3258 subclusters.

In the left-hand panels of Fig. 7, we see that δ_i for galaxies around both NGC 3268 and NGC 3258 are similar in magnitude, but when we set $\bar{v} = 2800 \text{ km s}^{-1}$, the galaxies centred around NGC 3268 have systematically lower δ_i , while those around NGC 3258 are larger, even though the elliptical galaxies have effectively the same velocity. The system of galaxies around NGC 3268 are better described by the velocity of the massive ellipticals, than

the mean velocity of the entire cluster. Meanwhile the system of galaxies around NGC 3258 are kinematically less consistent with the velocity of their central elliptical.

As mentioned above, four galaxies within 80 kpc projection of NGC 3258 and within its extended X-ray halo (Pedersen et al. 1997) are detected in $H\text{I}$. Three of these are blueshifted with respect to 2800 km s^{-1} , consistent with the rest of the NGC 3258 subcluster. We cannot rule out that the $H\text{I}$ population may be its own substructure seen in projection against NGC 3258, but we require redshift independent distances to disentangle the $H\text{I}$ and non- $H\text{I}$ detected galaxies. None the less, the locality of $H\text{I}$ in projection to NGC 3258, in addition to its weaker X-ray halo, is further evidence that the NGC 3258 subcluster is a younger system that has merged with the NGC 3268 cluster.

Lastly, while $\Delta \sim 1.4$ for the entire cluster for both choices of \bar{v} , if we consider only galaxies within the inset (bottom plots of Fig. 7): for $\bar{v} = 2747 \text{ km s}^{-1}$, $\Delta = 0.89$; while for $\bar{v} = 2800 \text{ km s}^{-1}$, $\Delta = 0.73$. Thus, it appears that galaxies on the outskirts of Antlia bias the calculation of the mean velocity of the cluster, and their δ_i values dominate the overall Δ_i , the collective impact of which is to wash out evidence for substructure in the centre.

5.1.2 The cluster outskirts

In addition to the NGC 3268 and NGC 3258 subclusters, we consider three more galaxy associations in Antlia. This is by no means a complete list, but we contemplate the most obvious collections of four or more galaxies which may have accreted on to the cluster together and not yet been disrupted by gravitational interactions. Based on the presence of $H\text{I}$ and/or star formation, we present them in increasing order of time since accretion. Alternatively, the ranking may be interpreted as one of increasing effect of the cluster environment on the evolution of the galaxies within the substructures:

(i) Subst 1, $(\alpha, \delta) = (157.3, -34.5)$: four galaxies, redshifted with respect to the cluster centre velocity, have high δ_i values and are at similar systemic velocities. All four of the galaxies are detected in $H\text{I}$ (Fig. 1) and the north-western most galaxy is one of the most $H\text{I}$ massive objects in the cluster (LEDA 082948). We postulate that collectively these make up the most recently arrived group to the cluster.

(ii) Subst 2, $(\alpha, \delta) = (157.6, -35.6)$: four galaxies which are closely associated in space, and vary little in δ_i . One of them is detected in $H\text{I}$ and one is detected in star formation. A fifth galaxy, seen in black, we consider unassociated.

(iii) Subst 3, $(\alpha, \delta) = (157.6, -35.2)$: seven galaxies in blue or black which behave similarly in δ_i when we change \bar{v} from 2711 to 2800 km s^{-1} . None of them are detected in $H\text{I}$ but one is detected by *WISE* to have star formation. These galaxies are quite close together, and stand out in δ_i from the nearby NGC 3268 subcluster.

Coincidentally, this list also ranks the substructure in decreasing projected distance to the cluster centre. During infall and assembly, the mixing of galaxy populations is incomplete, thus our proposed timeline is also consistent with a correlation between accretion time and distance to the cluster centre (De Lucia et al. 2012).

Subst 1 contains the most $H\text{I}$ bright galaxies and, based on their velocities, it could be either a foreground or background group falling into the cluster. The Hydra Cluster lies beyond the bounds of the KAT-7 mosaic to the north and east at higher redshifts (Radburn-Smith et al. 2006; Courtois et al. 2013), and thus the galaxies may be

in the background having come from a filament that connects Antlia with Hydra. In any case, we argue that the asymmetry in the velocity and spatial distribution of $H\text{I}$ detected galaxies (and substructure) is an intrinsic characteristic of accretion on to the Antlia Cluster at the present epoch.

Considering the overall spatial distribution of the Antlia population, we note that the line connecting NGC 3258 to NGC 3268, if extended to the north-east in Fig. 7, follows an overdensity of galaxies which have small δ_i . It is unclear whether this whole linear feature is related to the merger of NGC 3258 with NGC 3268, and/or connected with filamentary structures on larger scales. The small variation in δ_i would imply that much of the action in Antlia is occurring tangentially in the plane of the sky.

5.2 Dynamical mass estimates

We calculate the radial mass profile and estimate the total dynamical mass of the cluster based on all available redshift measurements which includes 124 galaxies with unique spectroscopic measurements. Nakazawa et al. (2000) use X-ray observations to estimate a total gravitating mass of $1.9 \times 10^{13} M_\odot$ internal to 250 kpc. By combining $H\text{I}$ and, in particular, 6dF redshifts we are able to probe out to a radius of 1 Mpc.

Mass estimates are commonly computed using the virial theorem, but can be biased. Heisler, Tremaine & Bahcall (1985) describe three alternative mass estimators for self-gravitating systems which provide consistent results, and work well even in systems where the number of galaxy members or available spectroscopic redshifts may be small, and the galaxies themselves have a range of masses. The four estimators are:

(a) virial mass,

$$M_{\text{VT}} = \frac{3\pi N}{2G} \frac{\sum_i V_{zi}^2}{\sum_i \sum_{i < j} 1/R_{\perp ij}} \quad (3)$$

(b) projected mass,

$$M_{\text{PM}} = \frac{f_{\text{PM}}}{G(N-1.5)} \sum_i V_{zi}^2 R_{\perp i} \quad (4)$$

$$f_{\text{PM}} = 32/\pi$$

(c) median mass,

$$M_{\text{Me}} = \frac{f_{\text{Me}}}{G} \text{med}_{i,j} (V_{zi} - V_{zj})^2 R_{\perp ij} \quad (5)$$

$$f_{\text{Me}} = 6.5$$

(d) average mass,

$$M_{\text{Av}} = \frac{2f_{\text{Av}}}{GN(N-1)} \sum_i \sum_{i < j} (V_{zi} - V_{zj})^2 R_{\perp ij} \quad (6)$$

$$f_{\text{Av}} = 2.8.$$

Fig. 8 shows how the mass estimators vary as we include galaxies at increasing projected radius. The cluster mass increases at a reasonable rate with projected radius, suggesting that the velocity selection criteria for cluster membership applied by previous authors reasonably defines the cluster, and that we have excluded outliers.

We use the average of these mass estimators to derive the cluster mass at various radii (Puche & Carignan 1988). Within 250 kpc, we find a cluster mass of $9.1 \pm 1.5 \times 10^{13} M_\odot$. This is 4.8 times greater than that suggested by X-ray observations within the same radius. Within 0.6–1 Mpc, the total mass of the cluster grows to

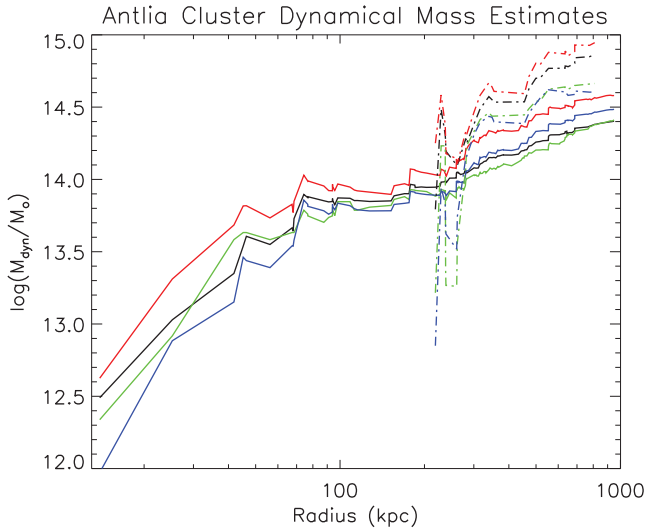


Figure 8. Dynamical mass estimates as a function of projected radius for the virial mass (black), projected mass (red), median mass (green), and average mass (blue). The solid lines are calculated using the full sample of 117 galaxies for which redshifts are available. Dot-dashed lines are calculated using only the 37 H I detected galaxies. These galaxies have a broader velocity dispersion and are not virialized, and therefore overestimate the dynamical mass of the cluster.

$2.6 \pm 0.6 \times 10^{14} M_{\odot}$. It may be interesting to note that virial mass and median mass track the closest with radius. According to Heisler et al. (1985), the median mass estimate is the least sensitive to outliers.

The mass estimates flatten at radii between 200 and 250 kpc. Combining the mass estimate and the velocity dispersion of galaxies within 250 kpc with the virial theorem, we calculate a cluster virial radius of 0.9 Mpc. Beyond 200–250 kpc, the mass estimators rise again out to 1 Mpc. Assuming the cluster population extends to 1 Mpc, the mass and velocity dispersion suggest a much larger virial radius of 3.5 Mpc. It is clear then that galaxies at greater distance from the cluster centre are less likely to be virialized.

Fig. 8 shows that the H I detected galaxies reside beyond at least 200 kpc from the cluster centre, and systematically overestimate the cluster mass for all estimators. Mass estimates which rely on the H I detected galaxies suggest a derived mass of $5.7 \pm 1.8 \times 10^{14} M_{\odot}$ and virial radius of 4.6 Mpc.

This picture is consistent with what has been found in Virgo, Fornax, and Coma: late-type galaxies have a broader velocity dispersion than early-type galaxies because they are free falling into a virialized system, and H I detected galaxies are on more radial orbits than those without (Wojtak & Łokas 2010). Beyond 250 kpc, there is a sizeable population of older cluster members which share the same volume as where the infall of H I objects is occurring from the surroundings.

5.3 NGC 3281: AGN and associated absorber

NGC 3281 is the brightest object detected in *WISE* and is projected at ~ 460 kpc from NGC 3268 (Fig. 9). The galaxy hosts one of the most nearby obscured, Compton thick AGNs ($N_{\text{H}} > 10^{24} \text{ cm}^{-2}$), and has been a target of X-ray and infrared observations to investigate the nature of the obscuring torus (Storchi-Bergmann, Wilson & Baldwin 1992; Simpson 1998; Vignali & Comastri 2002; Sales et al. 2011). The 3.4–4.6 μm colour is at the low end of the typical W1 –

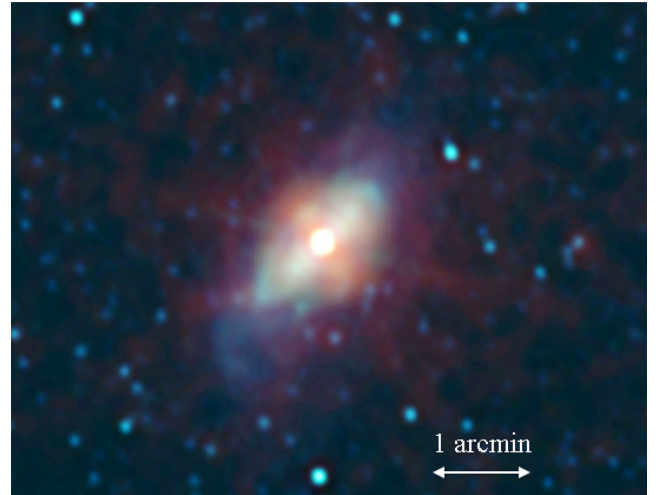


Figure 9. *WISE* mid-infrared view of NGC 3281. The image was constructed using the high-resolution sampling technique described in Jarrett et al. (2012), with colour combination W1 (3.4 μm) is blue, W2 (4.6 μm) is green, W3 (12 μm) is orange, and W4 (22 μm) is red.

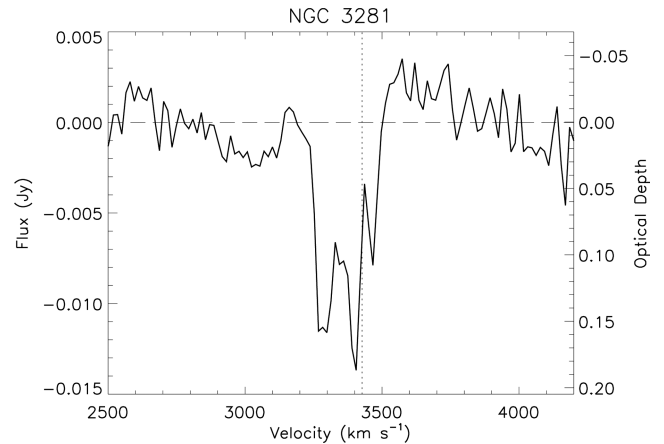


Figure 10. Absorption line spectrum of NGC 3281. The dotted line indicates the average redshift of results from the literature derived from optical emission lines.

W2 = 0.8 colour cut indicative of AGN activity (Jarrett et al. 2011; Stern et al. 2012), but its 6dF emission line spectrum indicate it is a Seyfert II, and Compton thick AGN are known to span a range of *WISE* colours (Fig. 4; Gandhi et al. 2014). Redshift measurements in the literature from optical emission lines give an average systemic velocity for the galaxy of $3427 \pm 26 \text{ km s}^{-1}$ (Martin 1976; Sandage 1978; Rubin et al. 1985; de Vaucouleurs et al. 1991; Jones et al. 2009).

We find no H I emission associated with the galaxy, but we detect it strongly in absorption (Fig. 10). The flux of the continuum source is $78.8 \text{ mJy beam}^{-1}$, and the peak of the absorption is $-13.4 \text{ mJy beam}^{-1}$ corresponding to a peak optical depth of 0.186. By integrating the optical depth over the width of the line and using $N_{\text{H I}} = 1.82 \times 10^{18} T_{\text{spin}} / f \int \tau(v) dv$ (Wolfe & Burbridge 1975), we calculate an H I column density of $5.05 \times 10^{19} T_{\text{spin}} / f \text{ cm}^{-2}$. The absorption spans 230 km s^{-1} and appears to be composed of three velocity components, two of which are blueshifted with respect to the systemic velocity of the galaxy and the third of which is redshifted.

The canonical value for T_{spin}/f is 100, but could be greater by a factor of 30 or more (Holt et al. 2006). Infrared observations give estimates for the dust temperature at the nucleus of NGC 3281 ranging from 300 to 1000 K or more (Sales et al. 2011; Winge et al. 2000). The 300 K estimate is extracted from a 130 kpc region around the nucleus and NGC 3281 is unresolved in radio continuum down to 180 kpc (Ulvestad & Wilson 1989). Assuming the dust couples well to the neutral gas, this would imply H I column densities of $N_{\text{H I}} = 1.5\text{--}5 \times 10^{22} \text{ cm}^{-2}$.

Controversially, an observation from the meridian transit Nançay radiotelescope claims a redshift of 3200 km s^{-1} for H I in emission (Theureau et al. 1998). However, Nançay is not an imaging telescope and the beam ($3.6 \times 22 \text{ arcmin}$) is significantly elongated. The determination of the instrumental baseline is likely the most uncertain aspect of accurately measuring the line strength and we speculate that the Nançay survey’s polynomial baseline fitting may have fit the absorption, overcorrecting, and mimicking emission in adjacent channels. This is speculation as they do not report a conscientious search for absorption line objects in the survey, but the Nançay spectrum (NED) shows a narrow trough consistent with where we see the strongest absorption in the KAT-7 mosaic.

Redshifted H I may be a source of fuel for the central AGN (e.g. van Gorkom et al. 1989; Morganti et al. 2009; Maccagni et al. 2014), while blueshifted gas may be the result of a jet driven outflow (e.g. Morganti, Tadhunter & Oosterloo 2005; Curran & Whiting 2010; Mahony et al. 2013). Storchi-Bergmann et al. (1992) adopt a systemic velocity of 3396 km s^{-1} and find evidence in NGC 3281 for an outflow in ionized gas from $[\text{O III}]$ and $\text{H } \beta$ emission lines. They argue that from the combination of -43 blueshifted and $+125 \text{ km s}^{-1}$ redshifted emission, the outflow is moving in a cone at 150 km s^{-1} with an inclination of 64° . Assuming a systemic velocity of $3427(3396) \text{ km s}^{-1}$, we find the blue-, redshifted absorption span $-174(-143)$, and $+55(+86) \text{ km s}^{-1}$, respectively. A detailed geometric analysis of NGC 3281 is outside the scope of this paper, but the optical emission lines are likely probing different size scales and different regions around the AGN than the H I absorption lines.

Alternatively, we cannot rule out that such an absorption line profile may be the result of H I distributed in a flattened rotating disc (Curran & Whiting 2010; Gereb, Morganti & Oosterloo 2014). A jet in the plane of the sky would produce both red- and blueshifted absorption on either side of the core (Peck, Taylor & Conway 1999).

6 SUMMARY AND CONCLUSIONS

The Antlia Cluster is dynamically one of the youngest clusters in the local Universe, consisting of two substructures in the process of merging as revealed by the galaxy populations, diffuse X-ray emission, and globular cluster populations surrounding the two central dominant elliptical galaxies, NGC 3268 and NGC 3258.

By combining KAT-7 H I spectral line commissioning observations with *WISE* broad-band infrared and existing optical spectroscopic redshifts, we present a significant advance in our understanding of the dynamical state of the cluster. From our H I mosaic, X-ray image, and kinematic DS test, we conclude that NGC 3268 is at the centre of an older, more dynamically mature structure, and nominally the core of the Antlia Cluster. Surprisingly, we find four H I detected galaxies in close projection to NGC 3258. However, even if they are a chance alignment, the kinematics suggest the system is secondary in the substructure hierarchy.

We identify the infall region beyond 200 kpc based on the spatial distribution and dynamical mass estimates from the H I detected galaxies compared to the predominantly early-type optical galaxy

population. Between 200 and 600 kpc radius, H I detected galaxies show signs that they have undergone significant stripping of their atomic gas. The largest cluster galaxies by stellar mass show the greatest H I deficiency, down by 2–3 orders of magnitude compared to their counterparts of the same stellar mass in the field. Beyond 600 kpc Antlia is still actively accreting galaxies from the surrounding environment, notably from the direction towards the Hydra Cluster.

Using the most complete sample of new and archival velocity redshifts, we estimate the cluster mass within a radius of 600 kpc to 1 Mpc to be $2.6 \pm 0.6 \times 10^{14} M_\odot$, and show that X-ray observations have underestimated the mass within the 200–250 kpc X-ray halo by nearly a factor of 5. We find that the velocity distribution of H I selected galaxies compared to the predominantly early-type sample are consistent with what has been modelled in other galaxy clusters: that late-type galaxies as traced by the H I content, are on primarily radial orbits, falling into a virialized system. Further, this accretion on to the Antlia Cluster is strongly asymmetric.

We use the DS test to identify infalling substructure within Antlia and on the outskirts by deviations from the global cluster kinematics. By examining galaxies within these substructures for their H I content and star formation rate, we propose a relative sequence for their accretion on to the Antlia Cluster which is also consistent with their distance from the cluster centre. We speculate that both the infall of individual galaxies from the field and the accretion of galaxy groups are important to the ongoing mass assembly of Antlia. Additionally, there is evidence that kinematics and asymmetry of the cluster H I distribution reflect the larger scale filaments and velocity flows of the supercluster environment towards the Hydra Cluster, although we are unable to say anything conclusive.

Finally, we report on the H I absorption associated with the Compton thick Seyfert II galaxy, NGC 3281. The object belongs to a group of galaxies we identify from the DS test, but it is of further interest for studies of both AGN outflows and star formation driven winds.

Despite the story of cluster assembly we have presented, we cannot rule out that some of the substructure we identify may be foreground or background objects seen in projection. Accurate distances require redshift independent measurements such as the Tully–Fisher relation. Finally, to gain a greater understanding of how the infalling population is impacted by the cluster environment requires high-resolution H I observations. These have the potential to reveal direct morphological evidence of interaction between the hot intracluster medium and cold ISM of the galaxies through ram pressure stripping.

Antlia is an important addition to the collection of clusters in the local Universe for which we are able to pursue detailed, high-resolution, multiwavelength synthesis of the stellar, gas, and dark matter content, which in turn can shed light on mass assembly, cluster formation, and galaxy transformation. MeerKAT will easily reach a factor of 10 deeper in H I mass and resolving both the gas morphology in H I and star formation in radio continuum.

Antlia’s apparent youth provides a local snapshot of what more mature clusters may have looked like in the past, and in greater detail than what we can directly observe in intermediate-redshift clusters. Antlia in combination with other massive systems have the potential to show whether the way in which galaxies are accreted on to clusters, in groups or individually, impacts the resulting galaxy cluster population. Although not the focus of this paper, the high fraction of S0 to elliptical galaxies in Antlia suggests that it may have a different formation history in contrast to other well-studied clusters. Whether Antlia contains clues as to the different formation mechanisms of Es versus S0s, remains to be seen.

ACKNOWLEDGEMENTS

We thank the anonymous referee for his/her helpful comments that improved the quality of this paper. We thank Tom Oosterloo for his insightful discussions and sharing with us his continuum subtraction code, and Thijs van der Hulst for his useful comments. The research of KH has been supported by a South African Research Chairs Initiative (SARChI), Square Kilometre Array South Africa (SKA SA) fellowship, and support from the European Research Council under the European Union's Seventh Framework Programme (FP/2007-2013)/ERC Grant Agreement nr. 291531. The work of TJ and CC is based upon research supported by the SARChI program of the Department of Science and Technology (DST), the SKA SA, and the National Research Foundation (NRF).

This work is based on observations obtained at the KAT-7 array which is operated by the SKA SA on behalf of the National Research Foundation of South Africa.

This publication makes use of data products from the *Wide-field Infrared Survey Explorer*, which is a joint project of the University of California, Los Angeles, and the Jet Propulsion Laboratory/California Institute of Technology, funded by the National Aeronautics and Space Administration (NASA).

This research has made use of the NASA/IPAC Extragalactic Database (NED) which is operated by the Jet Propulsion Laboratory, California Institute of Technology, under contract with the National Aeronautics and Space Administration.

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APPENDIX A: SUPPLEMENTARY FIGURES

Here, we provide the H_1 spectral line profiles for all the KAT-7 detections. All profiles have the same velocity scale, but the intensity is adjusted to the individual detections (Figs A1 and A2).

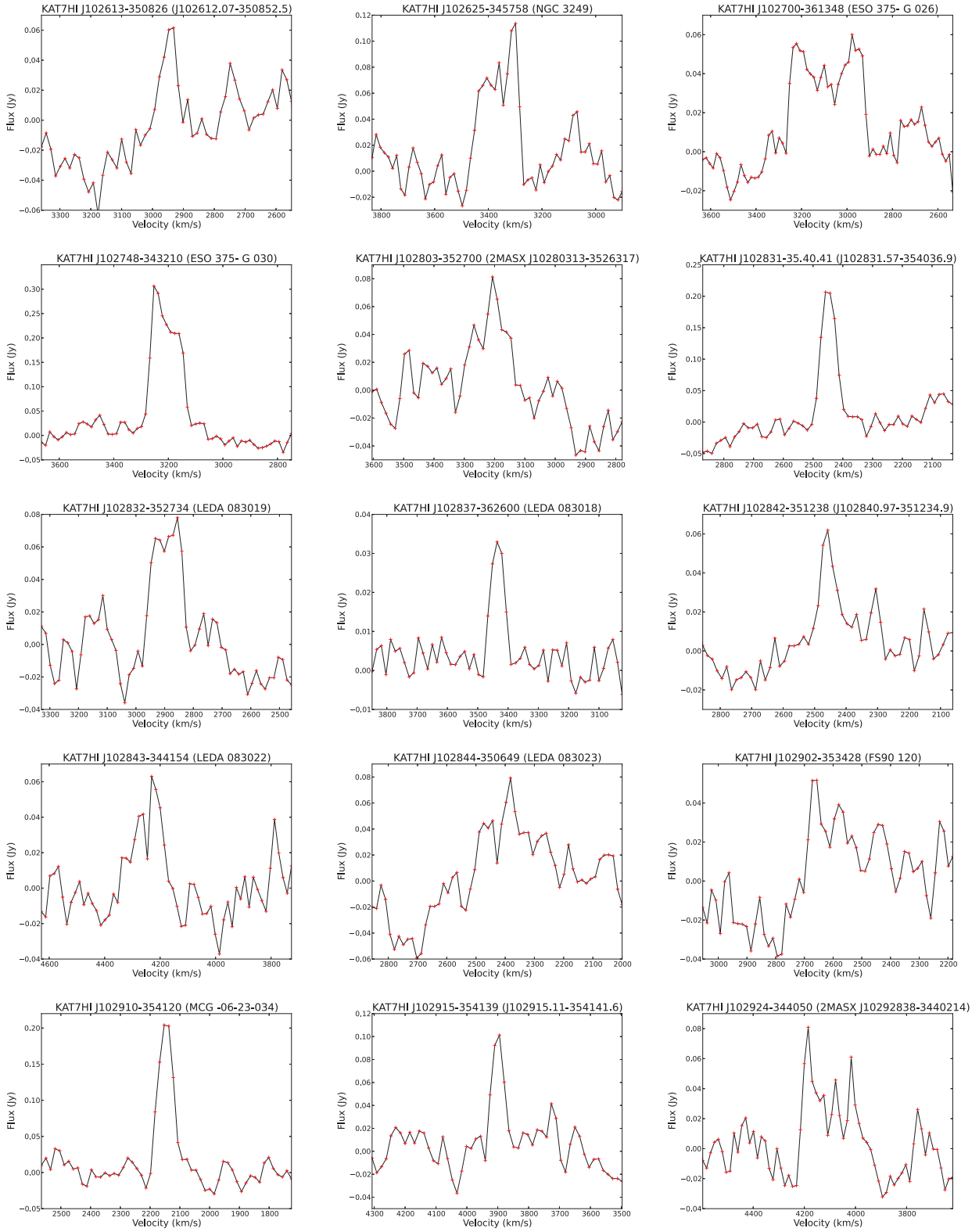


Figure A1. H I profiles of the KAT-7 detections detected in both the 15 and 31 km s⁻¹ cubes. Objects 1–30 appear in the same order as Table 2.

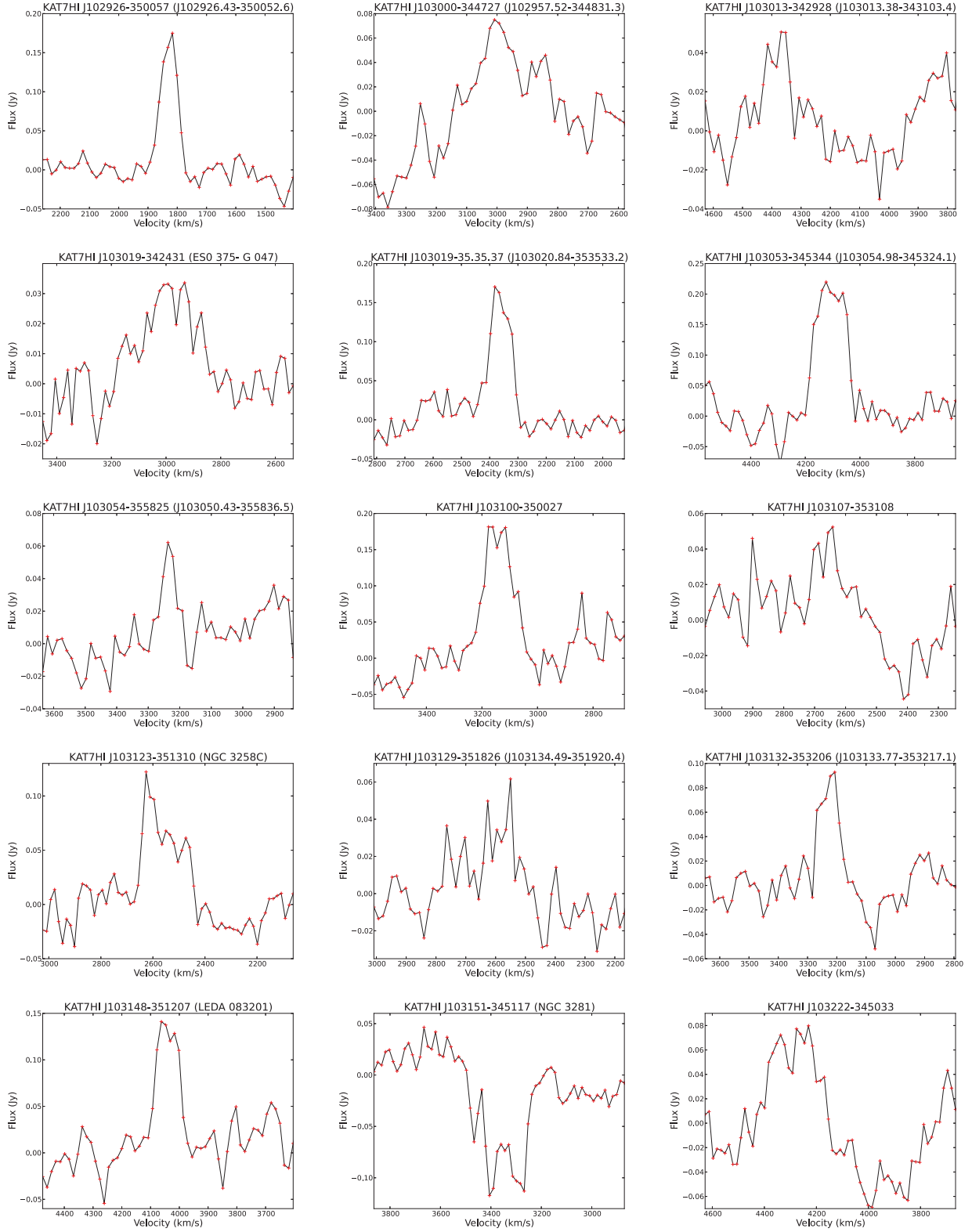


Figure A1 – continued

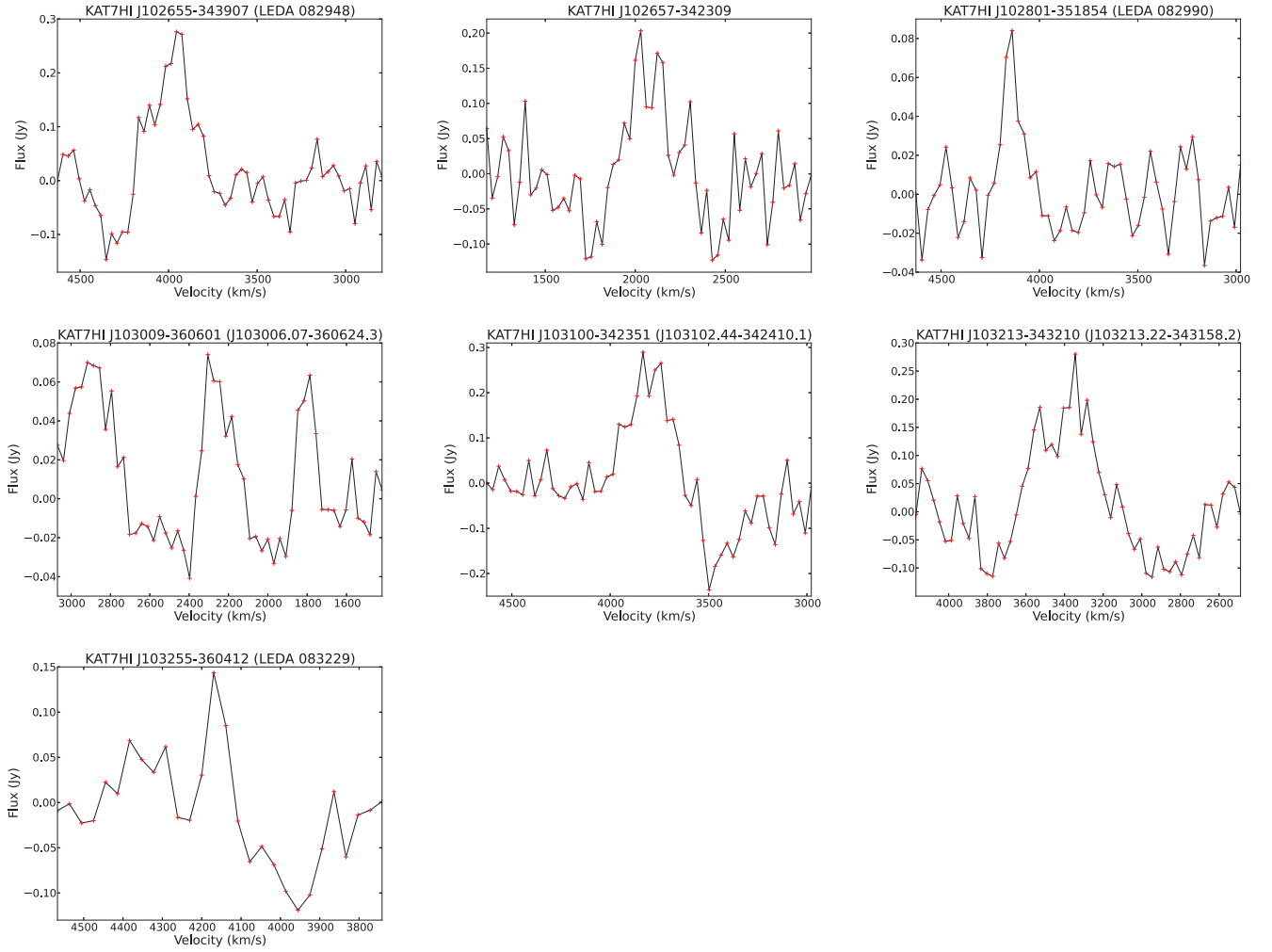


Figure A2. H I profiles of the KAT-7 detections which were only detected in the 31 km s^{-1} cube. Note, in all but the last panel the x-axis range is twice that in Fig. A1. Objects 31–37 appear in the same order as Table 2.

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